

COSMIC ORIGINS SPECTROGRAPH AND *FUSE* OBSERVATIONS OF $T \sim 10^5$ K GAS IN A NEARBY GALAXY FILAMENT¹

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ABSTRACT

We present a clear detection of a broad Ly α absorber (BLA) with a matching OVI line in the nearby universe. The BLA is detected at $z(\text{Ly}\alpha) = 0.01028$ in the high S/N spectrum of Mrk 290 obtained using the Cosmic Origins Spectrograph. The Ly α absorption has two components, with $b(\text{H I}) = 55 \pm 1 \text{ km s}^{-1}$ and $b(\text{H I}) = 33 \pm 1 \text{ km s}^{-1}$, separated in velocity by $v \sim 115 \text{ km s}^{-1}$. The OVI, detected by FUSE at $z(\text{OVI}) = 0.01027$, has a $b(\text{OVI}) = 29 \pm 3 \text{ km s}^{-1}$ and is kinematically well aligned with the broader H I component. The non-detection of other ions such as C II, Si II, Fe II, C III, Si III, C IV, Si IV and NV at the same velocity as the BLA and the OVI implies that the absorber is tracing highly ionized gas. The different line widths of the BLA and OVI suggest a temperature of $T = 1.4 \times 10^5 \text{ K}$ in the absorber. Photoionization, collisional ionization equilibrium as well as nonequilibrium collisional ionization models do not explain the ion ratios at this temperature. The observed line strength ratios and line widths favor an ionization scenario in which both ion-electron collisions and UV photons contribute to the ionization in the gas. Such a model requires a low-metallicity of ~ -1.7 dex, ionization parameter of $\log U \sim -1.4$, a large total hydrogen column density of $N(\text{H}) \sim 4 \times 10^{19} \text{ cm}^{-2}$, and a path length of $\sim 400 \text{ kpc}$. The line of sight to Mrk 290 intercepts at the redshift of the absorber, a megaparsec scale filamentary structure extending over $\sim 20 \text{ deg}$ in the sky, with several luminous galaxies distributed within $\sim 1.5 h^{-1} \text{ Mpc}$ projected distance from the absorber. The collisionally ionized gas phase of this absorber is most likely tracing a shock-heated gaseous structure, consistent with a few different scenarios for the origin including an over-dense region of the WHIM in the galaxy filament or highly ionized gas in the extended halo of one of the galaxies in the filament. In general, BLAs with metals provide an efficient means to study $T \sim 10^5 - 10^6 \text{ K}$ gas in galaxy halos and in the intergalactic medium. A substantial fraction of the baryons *missing* from the present universe is predicted to be in such environments in the form of highly ionized plasma.

1. INTRODUCTION

Cosmic baryon census estimates have shown that in our present universe only $\sim 6 - 10 \%$ of the primordial baryons are entrapped in galaxies. The bulk of the baryons have not collapsed into luminous structures but are in circumgalactic regions, and in the unvirialized large scale intergalactic filaments (Mulchaey et al. 1996; Fukugita et al. 1998; Fukugita & Peebles 2004). A significant fraction ($\sim 50\%$) of these baryons outside of galaxies has already been discovered by STIS/*HST* and *FUSE* UV spectroscopic observations of the $z < 0.5$ intergalactic medium (IGM). The low- z IGM exists as

a mix of multiple gas phases. The photoionized diffuse IGM with $T \lesssim 10^4 \text{ K}$ traced by the Ly α forest accounts for $\sim 30\%$ of the baryonic mass density (Shull et al. 1996; Penton et al. 2004; Lehner et al. 2006; Danforth & Shull 2008). The population of highly ionized metal line absorbers, notably OVI systems discovered closer to galaxies, possibly contain an additional $\sim 10\%$ of the mass-fraction if their ionization is dominated by collisions of electrons with ions in the plasma (Tripp et al. 2000b; Richter et al. 2004; Danforth & Shull 2005; Savage et al. 2005; Stocke et al. 2006; Danforth & Shull 2008; Tripp et al. 2008; Narayanan et al. 2010). The remaining $\sim 60\%$ of the baryonic mass-fraction in our present universe remains to be discovered. Hydrodynamic simulations of structure formation predict that this missing portion from the baryon inventory is in a gravitationally shock-heated phase with temperatures in the range of $T \sim 10^5 - 10^7 \text{ K}$ and densities of $n_{\text{H}} \sim (0.1 - 10) \times 10^{-5} \text{ cm}^{-3}$ (Cen & Ostriker 1999; Davé et al. 2001; Valageas et al. 2002). The simulations convey that the reservoir of cooler intergalactic gas gets heated to warm - hot temperatures as a result of shocks produced during the gravitationally collapse of matter to form galaxies and clusters in the over-dense regions of the cosmic web. Discovering this warm-hot intergalactic medium (WHIM) is one of the main science drivers for

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the *Cosmic Origins Spectrograph* (COS) on *HST* (Shull 2009).

Currently, UV absorption line spectroscopy is the most promising observational technique for detecting the component of the WHIM with $T \sim 10^5 - 10^6$ K. Among metal lines, OVI $\lambda\lambda 1031, 1037$ Å are the most sensitive for observations of high ionization gas, due to the high cosmic abundance of oxygen and the large values of oscillator strengths of the 1032 Å and 1038 Å resonant lines. Under collisional ionization equilibrium conditions, OVI ionization fraction reaches its peak value of $f_{\text{OVI}} = N(\text{OVI})/N(\text{O}) = 0.187$ at $T = 3.2 \times 10^5$ K (Gnat & Sternberg 2007). Thus, OVI can be a tracer of the *warm* phase (i.e. the lower temperature phase) of the shock-heated gas (Tripp et al. 2001; Savage et al. 2002; Danforth & Shull 2008). Theoretical estimations predict $\sim 20 - 30\%$ of the missing baryonic matter is traced by OVI absorption systems with $W_r(1032) \geq 20$ mÅ (Cen et al. 2001). However, observationally confirming the value of $\Omega_b(\text{OVI})$ has been a challenge. Detectable amounts of OVI can also be produced in low-density gas ($n_{\text{H}} \lesssim 10^{-5} \text{ cm}^{-3}$) photoionized and heated by a hard radiation field, in which case the absorber may not be tracing a large-scale structure (Savage et al. 1998; Prochaska et al. 2004; Thom & Chen 2008; Tripp et al. 2008; Oppenheimer & Davé 2009). The current sample of known OVI systems seems to suggest that they are a heterogeneous mix, with evidence for photons dominating the ionization in some cases, and electron collisions in others. Detailed ionization modeling with tight constraints on column densities, Doppler widths and velocities of H I, OVI and other metal ions are necessary to understand the ionization mechanism in each absorption system and its physical origin (Savage et al. 2005; Narayanan et al. 2010).

An alternative to using highly ionized metals is to search for absorption from the trace amounts of atomic hydrogen (ionization fraction, $f_{\text{H I}} = N(\text{H I})/N(\text{H}) < 10^{-5}$) present in the warm - hot gas. At $T \geq 10^5$ K, an absorber with a total hydrogen column density of $N(\text{H}) \sim 10^{19} \text{ cm}^{-2}$ will produce Ly α absorption with column densities of $N(\text{H I}) \lesssim 10^{14} \text{ cm}^{-2}$ and Doppler widths of $b(\text{H I}) \gtrsim 40 \text{ km s}^{-1}$. Accurate measurements of such broad lines in the quasar spectrum require high S/N , well-determined continua, and a good understanding of the fixed-pattern noise. Broad Ly α absorbers (BLAs) in the $z < 0.5$ universe were discovered in the existing archive of *HST*/STIS quasar spectra (Sembach et al. 2004; Richter et al. 2006; Lehner et al. 2007; Danforth et al. 2010). The most serious systematic uncertainty associated with these BLA measurements is whether the observed broad line width of Ly α is dominantly thermal. Due to the shallow and broad shape of the absorption profile, particularly when the H I column is low, it is often difficult to rule out, from moderate S/N spectra, the possibility of line blending or turbulence contributing to line width. High S/N spectroscopy is therefore crucial to confirm the thermally broadened Gaussian absorption profile characteristic of BLAs. At low- z BLAs account for a baryonic mass-density fraction similar to galaxies or possibly even greater (Richter et al. 2006; Danforth et al. 2010). Thus discovering more BLAs is

central to the continuing search for the reservoirs of missing baryons in the present universe. Observations of the low- z universe also opens up the possibility to search for an answer to the allied issue of how these baryons are distributed in relation to galaxies in large-scale structures.

In this paper, we report on a BLA in the very nearby universe detected in the high S/N spectra of Mrk 290 obtained with the *HST*/*Cosmic Origins Spectrograph*. The BLA is at the same velocity as an OVI absorber identified in the *FUSE* spectrum of the same sight-line. The BLA - OVI system is tracing collisionally ionized gas associated with a large-scale galaxy filament. We describe in detail the COS observations (sec 2.1), the properties of the absorber (sec 3), and the details on the galaxies in the foreground field (sec 4). The physical properties of the absorber are investigated using the predictions made by different ionization scenarios (sec 5). Based on the observed properties and the ionization analysis, we speculate on the origin of the absorption system and its baryonic content (sec 6). The cosmology that we adopt is Λ CDM with $\Omega_m = 0.3$, $\Omega_\Lambda = 0.7$ and a dimensionless Hubble parameter of $h = H_0/(100 \text{ km s}^{-1} \text{ Mpc}^{-1}) = 0.71$.

2. OBSERVATIONS

2.1. COS Observations

The design capabilities of *HST*/COS are described in detail by Green (2001); Froning & Green (2009) and in the updated COS Instrument Handbook (Dixon et al. 2010). The inflight performance of COS is discussed by Osterman et al. (2010) and in the numerous instrument science reports on the STScI COS website⁵. The COS observations of Mrk 290 were carried out on 28 October 2009 as part of cycle 17 COS GTO program ID 11541 (P.I. James Green). The details of the observations are listed in Table 1. The G130M and G160M COS gratings were chosen to obtain spectra in the far ultraviolet bands. The G130M observations were split into four exposures of ~ 964 second integrations at grating central wavelength positions of 1291 Å, 1300 Å, 1309 Å and 1318 Å. Similarly, the G160M observations were composed of four ~ 1200 second integrations with central wavelength settings of 1589 Å, 1600 Å, 1611 Å, and 1623 Å. All exposures were obtained in the time-tagged, photon-counting mode. The different grating central wavelength configurations resulted in photons of the same wavelength getting recorded at different regions of the detector which subsequently reduced the effect of fixed pattern noise in the final coadded spectrum. This set up also helped to cover the ~ 18 Å wavelength gap between segments A and B of the FUV detector for the various central wavelength settings.

The reduction of the raw frames into 1D spectra was carried out using the COS pipeline software Calcos (v2.11) with the appropriate calibration reference files. The one-dimensional science data files generated by the pipeline had a dispersion scale of 0.01 Å corresponding to 6 wavelength bins per resolution element. The individual integrations were coadded weighted by their respective

⁵ <http://www.stsci.edu/hst/cos/documents/isrs>

exposure times. The final spectrum had a wavelength coverage from 1136 Å to 1796 Å. The signal-to-noise ratios (per 18 km s⁻¹ resolution element) of this combined final spectrum at 1250 Å, 1450 Å, and 1650 Å are 35, 24, and 16 respectively.

Ghavamian et al. (2009) have determined the resolving power of the spectrograph by detailed modeling of the line spread function (LSF) at various wavelengths. The spectral resolution is found to be wavelength dependent with values in the range $\lambda/\Delta\lambda \sim 16,000 - 21,000$ for the G130M and G160M gratings, where $\Delta\lambda$ refers to the width at half strength of the LSF which has broad wings containing 25% of the LSF area. The resolution is maximum at near-UV wavelengths and declines monotonically towards lower wavelengths.

To determine zero-point offset in the COS spectrum we compare the velocity centroids of the $v_{\text{LSR}} = -138$ km s⁻¹ ISM absorption lines of Si II λ 1254 Å, Fe II λ 1608 Å and Al II λ 1671 Å from Complex C (Wakker 2006). These lines appear at $v_{\text{LSR}} = -159, -154$, and -153 km s⁻¹ respectively in the COS spectrum. A velocity correction of +17 km s⁻¹ was therefore applied to the COS data, which is approximately one resolution element. We estimate residual errors of ~ 5 km s⁻¹ (1 σ) in the velocity calibration.

2.2. FUSE Observations

The properties of the *FUSE* satellite and its performance are described in detail by Moos et al. (2000) and Sahnou et al. (2000). The processing of the *FUSE* data for this sight line is described in great detail elsewhere (Wakker et al. 2003; Wakker 2006). We only provide a short overview here. The archived *FUSE* observations for Mrk 290 were obtained over the time period from 2000 to 2007 (exposure IDs: P1072901, D0760101, D0760102, E0840101, E0840101, E0840102). The reduction of these data sets was carried out using version 2.4 of the *FUSE* calibration pipeline. The shifts in wavelength and central velocities in the individual reduced exposures were corrected by aligning the Milky Way ISM lines with the LSR velocity of the 21 cm H I absorption from Complex C ($v_{\text{LSR}} = -138$ km s⁻¹). Once the velocity shifts were applied to each observation, the data in the LiF1A and LiF2B channels of the detector were combined together to produce the final spectrum.

3. PROPERTIES OF THE ABSORBER

At a heliocentric velocity of $v_{\text{HELIO}} = 3081$ km s⁻¹, the $z(\text{OVI}) = 0.01027$ absorber is one of the nearest extragalactic OVI absorption systems. It was discovered by Wakker & Savage (2009) in the *FUSE* spectrum of Mrk 290 through the identification of OVI λ 1032 Å line. The detection was claimed as *uncertain* since the red member of the OVI doublet is confused with the Galactic Ar I λ 1048 Å line. A hint of the OVI λ 1038 Å came from the extra absorption wing seen at positive velocities of the Galactic Ar I line. The Ly α for this absorber was outside the *FUSE* spectral window, and the Ly β at this velocity suffers from Galactic C II contamination. Nonetheless, Wakker & Savage (2009) re-

ported the $> 3\sigma$ absorption detected at $\lambda \sim 1042.5$ Å as OVI λ 1032 Å after ruling out the possibility of the identified line being interstellar absorption or higher order Lyman lines along this sight line. The detection of Ly α in the COS G130M spectrum at $z(\text{Ly}\alpha) = 0.01028$ confirms the Wakker & Savage (2009) identification of this absorption system.

In Figures 1a - 1b we show the continuum normalized profiles of Ly α and OVI on a velocity scale centered at $z(\text{OVI}) = 0.01027$. Also displayed are regions of the spectrum where absorption from other prominent ionic species are expected. The Ly α profile is spread over a velocity interval of $\Delta v = 400$ km s⁻¹ in the rest frame of the absorber. The high S/N COS data show two distinct components in H I, separated in velocity by $v \sim 115$ km s⁻¹. Using the Fitzpatrick & Spitzer (1997) software, we fit Voigt profiles to the Ly α feature. The fit convolves the model profile with the G130M instrumental spread function at the observed wavelength of the Ly α feature. It is important that the specific line spread function determined for COS be used to fit the profile in order to minimize the impact of the non-gaussian wings on observed line profiles (see Ghavamian et al. 2009, for a detailed discussion).

The fit that best reproduces the observed absorption profile is a two component model with $\log [N(\text{H I}) \text{ cm}^{-2}] = 14.35 \pm 0.01$ dex and $b(\text{H I}) = 55 \pm 1$ km s⁻¹ at $v = 2 \pm 5$ km s⁻¹ and $\log [N(\text{H I}) \text{ cm}^{-2}] = 14.07 \pm 0.01$ dex and $b(\text{H I}) = 33 \pm 1$ km s⁻¹ at $v = 116 \pm 5$ km s⁻¹. This model fit is displayed in Figure 2. The fit reveals some weak residual absorption in the positive and negative velocity wings of the Ly α . A broader, shallower H I component could explain this excess absorption in the wings. However, because of blending with the stronger H I component at $v = 2$ km s⁻¹, the fitting routine is not able to yield meaningful parameters for an extra component at $v \sim -125$ km s⁻¹. The contribution to the total column density from such a shallow component is going to be marginal ($W_r \sim 30$ mÅ, $\log [N(\text{H I}) \text{ cm}^{-2}] \lesssim 13.0$) when compared with the contributions from the two stronger components at $v \sim 2$ km s⁻¹ and $v \sim 116$ km s⁻¹ respectively.

The $b = 55$ km s⁻¹ Doppler width for the $v = 2$ km s⁻¹ H I component suggests that this is a broad Ly α absorber (from here on, we often refer to this component in H I as the BLA). If the line is purely thermally broadened, then $b = 55$ km s⁻¹ line width implies a gas temperature of $T = 1.8 \times 10^5$ K. The most significant systematic uncertainty that could affect a b -measurement is blending of closely separated kinematic components. The residual absorption in the blue wing of the BLA suggests some kinematic complexity. However, the BLA and the $v \sim 116$ km s⁻¹ components of the Ly α profile appear highly symmetric suggesting very little influence of this low column density component on the measured b values.

We take note of the fact that the Ly α profile could be subjected to saturation that is unresolved by COS (FWHM ~ 17 km s⁻¹). The presence of instrumentally unresolved saturation becomes evident when appar-

ent column density profiles of different electronic transitions from the same energy level in an ion are compared (Savage & Sembach 1991). Such an analysis is not feasible in the case of the $z = 0.01027$ absorber as the Ly β covered by *FUSE* is strongly affected by contamination from Galactic ISM C II absorption. The Ly γ , which is ~ 18 times weaker compared to Ly α in line strength ($f\lambda$), falls at the edge ($\lambda \sim 982.5$ Å) of the *FUSE* coverage where the quality of the data is poor ($S/N \sim 4 - 6$ per resolution element). Even so, the H I upper limit on column density of $10^{15.1} \text{ cm}^{-2}$ given by the low S/N Ly γ data suggests that there is no significant saturation in either component of Ly α .

The COS spectrum offers coverage of several important low and high ions including C II, Si II, Fe II, Si III, C IV, Si IV and N V (see Figures 1a - 1b). The expected wavelengths for all these ions show non-detections at the 3σ significance level. We therefore quote in Table 2 upper limits on the column density for these nondetections. The lack of any low ionization absorption is indicative of high ionization conditions in the absorber. The C III $\lambda 977$ Å line is covered by *FUSE*, but falls at the edge of the spectrum where the S/N is poor (see Figure 1a). The line is not detected at $> 3 \sigma$ significance.

A Voigt profile fit to the OVI $\lambda 1032$ Å line yields a single component at $v = 1 \pm 5 \text{ km s}^{-1}$ with $\log [N(\text{OVI}) \text{ cm}^{-2}] = 13.80 \pm 0.05$ dex and $b(\text{OVI}) = 29 \pm 3 \text{ km s}^{-1}$. We use a Gaussian of FWHM = 25 km s^{-1} in this fit model to account for the *FUSE* instrumental smearing. The profile fit is displayed in figure 2. Using these fit parameters, we synthesize the OVI $\lambda 1038$ Å line and superimpose the profile on top of the data in figure 2. The comparison shows the uncontaminated positive velocity segment of the OVI $\lambda 1038$ Å absorption consistent with the OVI $\lambda 1032$ Å profile. We also use the apparent optical depth (AOD) technique of Sembach & Savage (1992), to derive $\log [N_a(\text{OVI}) \text{ cm}^{-2}] = 13.82 \pm 0.08$ dex and $b_a(\text{OVI}) = 35 \pm 8 \text{ km s}^{-1}$ for the OVI $\lambda 1032$ Å line. Our AOD measurements are similar to the Wakker & Savage (2009) measurement of the same feature (see Table 2). The b_a is the second moment of the apparent column density distribution, which when corrected for the *FUSE* instrumental spread gives a b -value that is consistent with the profile fit measurement.

In Figure 3, we plot the apparent column density profiles of Ly α and OVI. It is evident that the OVI absorption is kinematically well aligned with the BLA (i.e., the $v = 2 \text{ km s}^{-1}$ H I component). The formal velocity difference between the line centers is $\Delta v = v(\text{OVI}) - v(\text{Ly}\alpha) = -1 \pm 5 \text{ km s}^{-1}$. If the two ions are tracing the same gas phase, then the Doppler widths of $b(\text{H I}) = 55 \pm 1 \text{ km s}^{-1}$ and $b(\text{OVI}) = 29 \pm 1 \text{ km s}^{-1}$ implies a temperature of $T \sim 1.4 \times 10^5 \text{ K}$ with a non-thermal line broadening component of $b_{\text{nt}} = 26 \text{ km s}^{-1}$ ⁶. The derived temperature suggests that the $v = 2 \text{ km s}^{-1}$ H I component is by definition still a BLA (thermal

$b > 40 \text{ km s}^{-1}$), even though $\sim 50\%$ of the observed line broadening is non-thermal.

The wavelength region of the OVI $\lambda 1032$ corresponding to the narrower $v \sim 116 \text{ km s}^{-1}$ H I component is contaminated by a feature that is unrelated to this system. The wavelength region corresponding to the OVI $\lambda 1038$ line shows no OVI absorption associated with this H I component. The $b(\text{H I}) = 33 \pm 1 \text{ km s}^{-1}$ implies an upper limit on temperature of $T \leq 6.6 \times 10^4 \text{ K}$, which is consistent with photoionization. No other low or high ionization metal lines are detected at the velocity of this H I component. In Table 2, we separately list rest-frame equivalent width and column density upper limits for the various ions corresponding to the $v \sim 116 \text{ km s}^{-1}$ H I absorption.

4. ABSORBER - GALAXY ASSOCIATION

A search through the NED database reveals several extended sources within 1000 km s^{-1} of the BLA - OVI absorber in the foreground field of Mrk 290. The *left panel* of Figure 4a shows the projected distribution of these galaxies within a $\sim 16 h^{-1} \text{ Mpc}$ region. The distribution pattern suggests that the sight line is intercepting at the velocity $v_{\text{HELIO}} = 3081 \text{ km s}^{-1}$ of the absorber an over-dense filamentary structure in the cosmic web, with the galaxies organized around the axes of the filament. The megaparsec scale filamentary structure extends $\sim 15^\circ$ on the sky. In Figure 4a, we also plot the position of other AGN sight lines in the field. Among those, Mrk 876 lies 9.7° ($\sim 7 h^{-1} \text{ Mpc}$) further along the filament and shows Ly α , Ly β and OVI absorption at $v_{\text{HELIO}} = 3480 \text{ km s}^{-1}$ most likely associated with the filament (Wakker & Savage 2009). The PG 1626 + 554 sight line also probes this filament at a separation of $7^\circ.5$ ($\sim 5.6 h^{-1} \text{ Mpc}$) in the sky in the opposite direction to Mrk 876, but does not show any H I absorption near the expected velocity. The other AGNs shown in the Mrk 290 field (e.g. Mrk 279, Mrk 817, PG 1351 + 640) lie farther away from the filament and do not have any metal line or H I absorption detected near $v_{\text{HELIO}} \sim 3100 \text{ km s}^{-1}$.

The Mrk 290 field is inside the SDSS footprint. The galaxy sample in this region of this sky is nearly 100% complete down to a magnitude of $g = 17.5$ which corresponds to $\sim 0.03 L^*$ (Blanton et al. 2003). From the more than two dozen galaxies within a projected distance of $\sim 1.5 h^{-1} \text{ Mpc}$ from the absorber, NGC 5981, NGC 5982, NGC 5985, NGC 5987 and NGC 5989 are members of the loose group GH 158 (Geller & Huchra 1983). The galaxy distribution shows at least nine other galaxies within $\sim 1.5 h^{-1} \text{ Mpc}$ impact parameter and $\Delta v = 300 \text{ km s}^{-1}$ systemic velocity with respect to the absorber, but smaller in size and brightness compared to the GH 158 group members. Some of the other brighter, larger galaxies along this filament are also members of other galaxy groups (e.g. GH151).

From all these known extended foreground sources, the luminous galaxy NGC 5987 is closest in impact parameter to the absorber. NGC 5987 [$\alpha(\text{J2000}) = 15^h 39^m 57.4^s$, $\delta(\text{J2000}) = +58^\circ 04' 46''$] is at a projected distance of $\rho = 424 h^{-1} \text{ kpc}$ from the sight line. The galaxy has a redshift of $z = 0.0100$ ($v_{\text{HELIO}} =$

⁶ We solve for the temperature assuming that the observed b is a quadrature sum of the thermal and non-thermal b values, i.e., $b_{\text{obs}} = \sqrt{b_t^2 + b_{\text{nt}}^2}$.

3010 km s⁻¹ de Vaucouleurs et al. 1991) which gives it a $|\Delta v| = 71$ km s⁻¹ systemic velocity with respect to the BLA - OVI absorber. The galaxy has an estimated isophotal major diameter of $D_{25} = 51.7$ kpc, and is one of the largest spiral galaxies in the nearby universe (Romanishin 1983). Although the impact parameter is large, the small velocity separation and the large size of the galaxy makes it possible for the absorber to be in some way related to NGC 5987's circumgalactic environment. Such an inference is consistent with statistical results from a number of past absorber-galaxy surveys where $> 0.25 L^*$ galaxies were found within ~ 500 kpc of OVI absorbers (Tripp & Savage 2000a; Tripp et al. 2000b; Tumlinson et al. 2005; Stocke et al. 2006; Cooksey et al. 2008; Chen & Mulchaey 2009; Lehner et al. 2009; Wakker & Savage 2009) and also with results from cosmological simulations of galaxies and OVI absorbers (Ganguly et al. 2008).

NGC 5987 is a spiral galaxy of Sb morphology with a B -band magnitude of $B = 12.72$ or $M_B - 5 \log h = -20.4$ (de Vaucouleurs et al. 1991). This corresponds to a luminosity of $L \sim 2 L^*$ assuming $M_B^* = -19.6$. No published spectrum is available for this galaxy. In Figure 4b, we display the orientation of the galaxy with respect to the quasar sight line. The large impact parameter of the absorber with respect to the galaxy disk results in low surface flux of UV photons from the galaxy incident on the absorber. Thus, we do not expect the galaxy's radiation field to significantly influence the ionization of hydrogen or metals in the absorber (Fox et al. 2005; Narayanan et al. 2010). NGC 5987 also does not appear in the ROSAT all-sky survey bright source catalogue (Voges et al. 1999) and therefore is unlikely to harbor an active AGN that could photoionize the surrounding intergalactic gas.

Given the significant excess of galaxies with $v_{\text{HELIO}} \sim 3080$ km s⁻¹ in the field around Mrk 290, and the pencil beam size region probed by the BLA - OVI system ($v_{\text{HELIO}} = 3081$ km s⁻¹), it will be difficult to conclude whether the absorption is tracing the gaseous halo of one or the other galaxy in the filament, or an intergalactic gas cloud that is part of the large-scale filament. We return to this discussion in Sec 5.

5. IONIZATION CONDITIONS IN THE ABSORBER

Understanding the ionization, physical conditions and metallicity in the absorber is crucial to discriminate between the different possible explanations for its astrophysical origin. We explore a range of parameter space in both photoionization and collisional ionization models and assess whether the models are physically realistic. We focus our modeling on the $v \sim 2$ km s⁻¹ BLA component of the H I absorption which has OVI detected at the same velocity. No metals are detected for the other H I component at $v \sim 116$ km s⁻¹ and therefore useful constraints on the physical parameters cannot be obtained for this component from ionization modeling.

5.1. Photoionization

The $T = 1.4 \times 10^5$ K temperature implied by the different BLA and OVI line widths is much higher than the

kinetic temperature expected for a purely photoionized gas. Nonetheless, to elaborate on the difficulties associated with photoionization in explaining the origin of these lines, we investigate some best-fit photoionization models and their predictions. We use the photoionization code Cloudy [ver.C08.00, Ferland et al. (1998)] to solve for equilibrium models that reproduce $N(\text{OVI})$ and the limits on all the other ions. The source of ionization is assumed to be dominated by the extragalactic background radiation at $z = 0.01027$ whose shape and intensity is as modeled by (Haardt & Madau 2001). The UV background spectrum used has contributions from quasars and from the population of high- z young star forming galaxies. The photoionization models are calculated for different ionization parameters $\log U^7$, metallicity and the observed $N(\text{H I}) = 10^{14.35}$ cm⁻² in the broad H I component at $v = 1$ km s⁻¹. The $N(\text{OVI}) = 10^{13.80}$ cm⁻² and the limiting column density ratios of $N(\text{C IV})/N(\text{OVI})$, $N(\text{N V})/N(\text{OVI})$, $N(\text{Si IV})/N(\text{OVI})$ and $N(\text{C III})/N(\text{OVI})$ serve as additional constraints for the models. In the models, we assume the revised solar elemental abundance ratios of $[\text{C}/\text{H}]_{\odot} = -3.57$ dex, $[\text{O}/\text{H}]_{\odot} = -3.31$ dex, $[\text{N}/\text{H}]_{\odot} = -4.17$ dex and $[\text{Si}/\text{H}]_{\odot} = -4.49$ dex given by Asplund et al. (2009).

Shown in Figure 5 are the photoionization model predictions for column densities of the various ionic species, as a function of the ionization parameter. The observed $N(\text{C IV})/N(\text{OVI})$ and $N(\text{N V})/N(\text{OVI})$ limiting ratios restrict the ionization parameter to $\log U \gtrsim -0.6$, corresponding to $n_{\text{H}} \lesssim 6 \times 10^{-6}$ cm⁻³ and metallicity to -1.8 dex of solar. The $N(\text{H I}) = 10^{14.35}$ cm⁻² and the $\log U$ dependent ionization fraction of $f_{\text{H I}} \lesssim 10^{-5}$ implies an absorbing region with a total baryonic column density of $N(\text{H}) \geq 4 \times 10^{19}$ cm⁻². The very low density and high total hydrogen column density results in a line of sight thickness of $L \gtrsim 3$ Mpc for the absorbing structure. If the absorber is unvirialized, the large path length would result in lines with $v > H(z) L \sim 210$ km s⁻¹ width due to Hubble flow, which is inconsistent with the observed H I and OVI velocity widths. If the absorbing structure is well within the virialized halo of NGC 5987, it would have decoupled from cosmic expansion. The kinematic properties of the BLA would then largely depend only on the internal velocity field along the line of sight and the rotational velocity of the galaxy halo. It would be remarkable for a megaparsec scale absorbing region to not have any velocity field capable of inducing asymmetry or kinematic complexity in the line profiles. To the contrary, we find no such evidence in the observed line profiles of the BLA and OVI which appear kinematically simple, symmetric and Gaussian. The predictions of the models are further complicated by the fact that the photoionization equilibrium temperature of $\sim 0.6 \times 10^5$ K is significantly smaller than the $T = 1.4 \times 10^5$ K given by the different BLA and OVI line widths. Thus photoionization does not appear as a dominant process controlling the ionization state. We therefore turn to the alternate scenario of collisional ionization.

⁷ Ionization parameter is defined as the ratio of the number density of photons with $E \geq 13.6$ eV to the hydrogen density, $U = n_{\gamma}/n_{\text{H}}$

5.2. Collisional Ionization

5.2.1. Collision Ionization Equilibrium Models

The temperature suggested by the observed line widths is consistent with collisional ionization conditions. First we consider the simplest scenario of collisional ionization equilibrium (CIE) in which at a constant kinetic temperature the fraction of ions remain time-independent. In Figure 6, we compare observations with column density predictions from CIE models of (Gnat & Sternberg 2007) for a metallicity of 1/10 solar. We find that CIE models fail to explain all observed high ion column density limits simultaneously for acceptable temperatures. The predicted $N(\text{OVI})$ is greater than $N(\text{C IV})$ and $N(\text{Nv})$ for $T \gtrsim 2 \times 10^5$ K, which is higher than the temperature given by the line widths of the BLA and OVI. Between the temperature range of $T = 3.2 \times 10^5$ K, where the OVI ionization fraction peaks and $T = 10^5$ K, the OVI fraction declines rapidly by a factor of ~ 4 dex. Thus, at $T = 1.4 \times 10^5$ K, $0.1Z_\odot$ metallicity and $N(\text{H I}) = 10^{14.35} \text{ cm}^{-2}$, the CIE predicted $N(\text{OVI})$ is a factor of ~ 3 dex smaller than the measured value and the $N(\text{C IV})$ and $N(\text{Nv})$ are significantly overproduced for solar relative elemental abundances.

5.2.2. Hybrid of CIE & Photoionization Models

Even when conditions are favorable for collisional processes to dominate the ionization state, a realistic modeling of absorption systems should take into account the effect of UV metagalactic flux from quasars and star forming galaxies on the absorber (Danforth et al. 2006; Richter et al. 2006; Tripp et al. 2008; Danforth et al. 2010). We therefore compute a series of hybrid models using Cloudy, allowing both collisional and photoionization reactions to contribute to ion fractions. The temperature in Cloudy is set to $T = 1.4 \times 10^5$ K given by the different BLA - OVI line widths. The model that Cloudy would converge on will have the ionization and recombination in equilibrium even though the heating and cooling rates are out of balance, due to the fixed temperature. In Figure 7 we show the column density predictions for these hybrid models at $T = 1.4 \times 10^5$ K and different $\log U$. The density of the absorber in these models is determined by the input number density of hydrogen ionizing photons and the ionization parameter. The hybrid models that reproduce the measured $N(\text{OVI})$ and the column density ratios of OVI with C IV and Nv require $[Z/H] \sim -1.9$ dex, $\log U \sim -0.6$, $n_{\text{H}} \sim 4 \times 10^{-6} \text{ cm}^{-3}$, $N(\text{H}) \sim 1.8 \times 10^{20} \text{ cm}^{-2}$ and thus absorber path lengths $L \sim 12$ Mpc. The expected line width from Hubble broadening for such a large path length is incompatible with the observed H I and OVI line widths and therefore a CIE - photoionization hybrid model can be rejected.

5.2.3. Non-equilibrium Collisional Ionization Models

The temperature $T \sim 10^5$ K implied by the data is in the regime where radiative cooling is efficient. The presence of metals can enhance the rate of cooling of gas such that significant departures of ion fractions from CIE will happen at these temperatures. The time dependent ionization fractions in a gas radiatively cooling from an

equilibrium temperature of $T \sim 10^6$ K were calculated by Gnat & Sternberg (2007). In Figure 8, we illustrate the non-equilibrium collisional ionization (noneq-CI) curves for such a cooling gas flow using their model for metallicities of Z_\odot and $0.1Z_\odot$. In non-equilibrium cooling models, the recombination process lags behind the declining temperature such that the ionization fraction for high ions will be higher compared to CIE for temperatures in the range $10^{4.5} - 10^{5.3}$ K. The effect is pronounced for higher gas metallicities. At solar metallicity and $T = 1.4 \times 10^5$ K, the OVI ion fraction in noneq-CI models is $f(\text{OVI}) = 2.63 \times 10^{-3}$, a factor of 1.4 dex larger than the OVI ion fraction at CIE. From Figure 8, it can be inferred that for $Z \sim 0.3 Z_\odot$, the $N(\text{OVI})$ given by the model will be in agreement with the observed value for $T = 1.4 \times 10^5$ K. However, for the same model, the predicted $N(\text{C IV}) \sim N(\text{OVI})$ whereas observations suggest $N(\text{OVI})/N(\text{C IV}) \gtrsim 8$. The C IV prediction from this model would be consistent with the data only if the carbon to oxygen abundance ratio is reduced from solar by ~ 1 dex. Thus, a simple noneq-CI scenario may not explain in a satisfactory manner the ionization in this absorber. However by including in these noneq-CI models the influence of UV photons from the extragalactic background radiation, an acceptable ionization solution can be obtained. This is elaborated in the following subsection.

5.2.4. Hybrid of noneq-CI & Photoionization Models

In Figure 9, we display the changes to OVI and C IV fractions when collisions of ions with both electrons and photons are simultaneously influencing the ionization in the gas. The models in Figure 9 were computed using Cloudy for a fixed temperature of $T = 1.4 \times 10^5$ K given by the BLA - OVI line widths and a fixed intensity of the Haardt & Madau (2001) (QSOs and galaxies) extragalactic radiation field determined for the redshift of the absorber. The models are therefore a hybrid of equilibrium photoionization and equilibrium collisional ionization reactions. The changes perceived for the ion ratios in these hybrid models can be used to estimate how the ion ratios would be affected in a noneq-CI cooling gas scenario when photoionization is included. It should not be forgotten that several broad assumptions are inherent in these ionization models (plane-parallel slabs of uniformly cooling gas with no internal temperature gradient). Few real absorbers would resemble closely with these conditions. We should only expect these computational models to converge on a general acceptable solution rather than an exact fit to the full range of observed ionic column densities. Our correction allowing for the effects of photoionization on the noneq-CI cooling gas is only approximate. A proper calculation would need to follow the changing ionizing conditions in gas cooling from $T \sim 10^6$ K to $T \sim 1.4 \times 10^5$ K when simultaneously subjected to the extragalactic radiation field.

From the hybrid models shown in the *left panel* of Figure 9, we draw the following conclusions. For low values of $\log U$ (i.e. high density), ionization by photons does not contribute significantly to the OVI to C IV ratio. For $\log U < -4$, $N(\text{OVI})/N(\text{C IV})$ is similar to its CIE value of $N(\text{OVI})/N(\text{C IV}) \sim -3.4$ dex at $T = 1.4 \times 10^5$ K. With

an increase in ionization parameter (i.e., a large number of ionizing photons per electron), photoionization begins to increasingly influence the ionization fraction of OVI such that over the interval $-4.5 \leq \log U \leq -1.5$, the $N(\text{OVI})/N(\text{H I})$ and $N(\text{OVI})/N(\text{C IV})$ ratios are raised by a factor of ~ 3 dex. This region of the parameter space corresponds to the interval that separates models where OVI is predominantly collisionally ionized from models where photoionization begins to predominate. Within this $\log U$ interval the C IV ionization fraction remains the same as CIE and begins to decline at $\log U > -2.5$. We find that the change to the $N(\text{OVI})/N(\text{C IV})$ ion column density ratios seen in Figure 9 (*left panel*) is not sensitive to metallicity. This increase in the ionization fraction of OVI from pure collisional values, due to the introduction of the ionizing radiation field can be used to interpret hybrid models of noneq-CI and photoionization.

For $[Z/H] = -2.0$, the noneq-CI model predictions at $T = 1.4 \times 10^5$ K are $N(\text{OVI}) \sim 10^{10.6} \text{ cm}^{-2}$, $N(\text{C IV}) \sim 10^{13} \text{ cm}^{-2}$, $N(\text{NV}) \sim 10^{12.3} \text{ cm}^{-2}$ (refer *right panel* Figure 9). The NV prediction is already consistent with the upper limit given by the data, whereas C IV is overproduced by a small amount (0.1 dex higher), and OVI underproduced by ~ 3.2 dex. In the hybrid model shown in the *left panel* of Figure 9, we can see that the OVI ionization fraction is enhanced from noneq-CI by ~ 3.2 dex for $\log U \sim -1$, whereas the C IV ionization fraction is diminished by ~ 0.5 dex, making both column density predictions consistent with the measurements. However, the density $n_{\text{H}} \sim 10^{-5} \text{ cm}^{-3}$, and total hydrogen column density of $N(\text{H}) \sim 6 \times 10^{19} \text{ cm}^{-2}$ corresponding to this ionization parameter implies a line of sight of ~ 2 Mpc through the absorbing region. Such a large path length is inconsistent with the observed Ly α and OVI line widths. Lowering the ionization parameter would decrease the amount of OVI and increase the amount of C IV making the predictions disagree with measurements.

For the *hybrid* model to agree with measurements, the $\log U$ has to be reduced (thereby narrowing the path length) simultaneously raising the metallicity in the collisionally ionized phase. For $[Z/H] = -1.6$, the $N(\text{OVI})$ will increase by ~ 0.7 dex from its $[Z/H] = -2.0$ value (see Figure 9) whereas the $N(\text{C IV})$ will increase by ~ 0.4 dex. In the presence of ionizing photons the predicted $N(\text{OVI})$ will increase by an additional 2.5 dex making it equal to the measured value for $\log U \sim -1.5$. For this *hybrid* model, the $N(\text{C IV})$ will be ~ 0.4 dex smaller than the noneq-CI predicted value of $N(\text{C IV}) \sim 10^{13.4} \text{ cm}^{-2}$ making it approximately consistent with the 3σ upper limit of $10^{12.9} \text{ cm}^{-2}$. The $\log U \sim -1.5$ corresponds to $n_{\text{H}} \sim 4 \times 10^{-5} \text{ cm}^{-3}$ and a total hydrogen column density of $N(\text{H}) \sim 4 \times 10^{19} \text{ cm}^{-2}$ implying a path length of ~ 320 kpc for the absorbing region. Such a path length is consistent with the Ly α and OVI line widths. The Gnat & Sternberg (2007) noneq-CI models (shown in Figures 8 and 9) are available only for values of metallicity that differ by 1 dex. We have therefore used approximate scaling estimates for the noneq-CI column densities at $[Z/H] = -1.6$.

There is a narrow range in metallicity and ionization

parameter space where the *hybrid* model predictions will be consistent with the measured column densities and line widths. The values of $[Z/H] = -1.6$ and $\log U \sim -1.5$ are an upper limit on the metallicity and a lower limit on the ionization parameter, respectively. Increasing the metallicity would result in C IV being overproduced from the OVI phase. For a lower metallicity of $[Z/H] = -1.8$, the $\log U \sim -1.3$ in order to recover the observed $N(\text{OVI})$. This *hybrid* model would also be consistent with the $N(\text{C IV})$ upper limit. However, the suggested path length is ~ 500 kpc. Thus $[Z/H] = -1.8$ and $\log U \sim -1.3$ are therefore lower and upper limits respectively. Thus, within the framework of this simple noneq-CI and photoionization *hybrid* model, the BLA - OVI gas phase has a metallicity of $[Z/H] = -1.7 \pm 0.1$ dex and ionization parameter $\log U = -1.4 \pm 0.1$ dex. The implied total hydrogen column density is very large ($N(\text{H}) \sim 4 \times 10^{19} \text{ cm}^{-2}$) and therefore interesting. It is intriguing that even though the BLA is tracing collisionally ionized hydrogen at $T \sim 1.4 \times 10^5$ K, in the *hybrid* model much of the OVI is produced through photoionization.

It is important to emphasize here that the estimated metallicity and ionization parameter are dependent on the simplistic assumptions intrinsic to the ionization models and hence should not be considered as very robust. For example, the predicted $\sim 2\%$ solar metallicity and $n_{\text{H}} \sim 4 \times 10^{-5} \text{ cm}^{-3}$ density would imply a radiative cooling time-scale of ~ 25 Gyr for the gas to cool from an initial temperature of $T \sim 10^6$ K⁸. However, this time-scale estimation depends on the density and metallicity of the gas, and the initial temperature from which the gas begins cooling, all of which are not well determined for this absorber. A lower ionization parameter (higher density), higher metallicity or lower initial temperature would all result in lower cooling time-scales values that are consistent with the Hubble time-scale.

The above *hybrid* model also has the limitation that it does not take into account the enhanced metal line cooling due to the increase in line emissions from excess OVI produced by photoionization. A ~ 3 dex increase in the OVI column density could lower the cooling time scales thereby altering the ionization states of C IV and OVI. A more accurate treatment of the *hybrid* model should incorporate photoionization by external radiation into the Gnat & Sternberg (2007) ionization calculations and subsequently follow the ion fractions as the gas cools from $T \sim 10^6$ K. This is a computationally elaborate task and is beyond the scope of this paper.

6. ASTROPHYSICAL ORIGIN OF THE BLA - OVI ABSORBER

In contemplating the nature of this BLA - OVI absorber, we are guided by two key insights, namely (1) the absorption is produced by gas at $T = 1.4 \times 10^5$ K whose ionization is likely governed by both collisional non-

⁸ The radiative cooling time for a low-density plasma is given by $t_{\text{cool}} = 1.5kT/(n_{\text{H}}\Lambda(T))$ (Sutherland & Dopita 1993), where k is the Boltzmann's constant and $\Lambda(T)$ is the temperature dependent radiative cooling efficiency of plasma. For gas at $T = 10^6$ K with $\sim 2\%$ solar metallicity, $\Lambda \sim 10^{-24} \text{ erg s}^{-1} \text{ cm}^3$ (Gnat & Sternberg 2007)

equilibrium and photoionization reactions (2) the absorbing structure is part of a large-scale filament of galaxies with several luminous galaxies within a $\sim 1.5 h^{-1}$ Mpc projected distance. In view of these, a few different scenarios are possible for the origin of the BLA - OVI absorption. The information obtained from probing a pencil beam region of the absorber may not be sufficient to discriminate between these different physical scenarios. Given the uncertainties encountered in the ionization models, we treat the inferences from the models as secondary to what the data is directly suggesting.

Starburst driven outflows with large velocities from galaxies can shock-heat and increase the ionization of the surrounding gas to produce OVI (Shull & McKee 1979; Heckman et al. 2001; Indebetouw & Shull 2004; Oppenheimer & Davé 2009). Observationally such a physical scenario is supported by the detections of OVI systems with $[O/H] \geq 0.2$ dex in the Galactic halo as well as the extended halos of other galaxies within ~ 500 kpc (Savage et al. 2002; Fox et al. 2004; Tumlinson et al. 2005; Stocke et al. 2006; Wakker & Savage 2009). However, gas associated with chemical feedback from star-formation typically have metallicity higher than the ~ -1.7 dex given by the noneq-CI and photoionization *hybrid* models for the BLA - OVI absorber.

A shock-heated origin for OVI was suggested by Stocke et al. (2006). In their sample of OVI absorbers and galaxies, 43% (16/37) of the absorbers were associated with pairs of Ly α clouds within $\Delta v = 50 - 200$ km s $^{-1}$ of each other. Both the OVI and the Ly α absorptions $[N(H\text{ I}) \geq 10^{13.2} \text{ cm}^{-2}]$ were also within 500 kpc of L^* and 100 kpc of $0.1 L^*$ galaxies. About half of these OVI in Ly α pairs were also found associated with galaxy groups. The velocity differences between the Ly α clouds were sufficient to create post-shock temperatures for OVI to be produced. The $z = 0.1212$ OVI absorber towards H 1821 + 643 is another example in which the hydrogen is seen in a cluster of Ly α absorbers. In this example, the various Ly α components were within ~ 100 km s $^{-1}$ of each other (Tripp et al. 2001). The H I component kinematically aligned with the OVI in that absorber was a BLA with $b \sim 85$ km s $^{-1}$ (see Table 3). Several galaxies with impact parameters ranging from 144 kpc to 3.1 Mpc were also identified within $|\Delta v| \sim 500$ km s $^{-1}$ of the absorber implying that the sight line was intercepting a galaxy group (Tripp et al. 1998, 2000b). Collisional ionization was inferred as the likely mechanism for the ionization in the OVI - BLA gas phase in that system (Tripp et al. 2001). The $z = 0.01027$ OVI absorber is also associated with a pair of Ly α absorbers separated by $|\Delta v| \sim 115$ km s $^{-1}$. Gas clouds moving at those relative velocities, colliding with one another, can shock-heat gas to temperatures of $T_s = 1.38 \times 10^5 (v_s/100)^2 \sim 1.8 \times 10^5$ K (assuming $\gamma = 5/3$ Draine & McKee 1993), consistent with the temperature for the BLA - OVI gas phase. Shocks can therefore be one of the mechanisms for the ionization in the absorber.

In a study of OVI absorbers and galaxies along select quasar fields, Chen & Mulchaey (2009) find some OVI absorbing galaxies with disturbed morphologies indicating past events of mergers or interactions. They argue

that such galaxies likely have a higher cross-section for OVI absorption due to tidally stripped gas contributing to the halo gas content. The BLA - OVI absorber is also in a galaxy excess environment (~ 30 galaxies within $1.5 h^{-1}$ Mpc projected distance) where the probability of galaxy pair interactions is high. Consequently, it is plausible that the absorption is from a shock-heated fragment of gas tidally stripped by NGC 5987 from one of the smaller galaxies in the large-scale filament. The oxygen abundance in the ISM of low-mass irregular galaxies is found to be low ($< 0.1 Z_\odot$, Skillman et al. 1989; Kennicutt & Skillman 2001). The tidally stripped interstellar gas of such a dwarf satellite would be consistent with the low-metallicity for the BLA - OVI absorber given by the *hybrid* models. The possible difficulty associated with this scenario is that tidally stripped gas is likely to produce absorption that is kinematically complex (Gibson et al. 2000; Wakker et al. 2003; Fox et al. 2010), which is not seen in the BLA and OVI profiles.

The BLA - OVI absorption can also be associated with the warm-hot intergalactic gas in the galaxy filament. As the dominant phase of baryons at $z \sim 0$, the WHIM is expected to be the primary reservoir of gas feeding mass into galaxies. The $T \gtrsim 10^5$ K suggested by the data for the absorber corresponds to the temperature predicted for the moderately over-dense cooler component of the WHIM (Kereš & Hernquist 2009). The BLA - OVI absorber thus presents a unique opportunity to probe the interaction of WHIM tied-in with galaxy halos. The $\sim 2\%$ solar metallicity for the absorber, given by the *hybrid* model, is consistent with the range measured for the low- z IGM (e.g., Danforth & Shull 2005). Even if the sight line to Mrk 290 is intercepting a nearby galaxy's (e.g., NGC 5987) extended halo, the absorption could still be from a shock-heated gas cloud accreted from the nearby WHIM. Such inflows of matter from WHIM are expected in the cold gas accretion models of galaxy formation (Kereš et al. 2005; Brooks et al. 2009; Kereš & Hernquist 2009). The presence of paired Ly α absorbers, and an origin in a galaxy filament, makes the absorber an excellent candidate for shock-heated WHIM (Stocke et al. 2006).

7. BLAS WITH METALS AS A TRACER OF WHIM

In this section we remark on why BLAs with metals are an important class of absorption system. Considerable observational effort has been invested to detect warm - hot collisionally ionized gas in the low- z universe, as it is predicted to be the dominant reservoir phase of ordinary matter (e.g., Bregman 2007). The broad-Ly α absorbers (by definition thermal $b(H\text{ I}) > 40$ km s $^{-1}$) offer a metallicity independent means to trace warm - hot gas (Richter et al. 2006; Danforth et al. 2010). However the biggest uncertainty associated with the temperature measurement in BLAs is in ascertaining the fraction of line broadening that is thermal. Blending of closely separated absorption components and turbulence within the gas can significantly contribute to the broadening of the absorption. The extent of this contribution is not readily assessable from the broad H I profile itself. Detecting metals associated with BLAs helps to resolve this problem. From the combined b -values for the metal line and

H I, one can solve for the thermal and non-thermal contributions to the line width and establish a tight constraint on the temperature of the gas. The temperature measurement is essential to determine with a certain level of precision the ionization fraction of H I and the baryon content in BLAs. This value of $N(\text{H})$ can be used in conjunction with single-phase solution to also get an approximate measure of the abundance of metals in the gas. The absorption system discussed in this paper is an example in which such an analysis was possible.

Metals have been detected in other BLAs. In table 3 we provide a partial list of BLAs with associated metal line detections with $|v(\text{BLA}) - v(\text{metal})| \leq 10 \text{ km s}^{-1}$. The close alignment between the H I and metal absorption (usually OVI) can be evidence for the origin of both species in a single gas phase. For each example listed in table 3, the temperature derived from the $b(\text{H I})$ and $b(\text{OVI})$ values is $T \geq 10^5 \text{ K}$, suggesting a dominant thermal broadening component for the line widths. The ionization in the gas phases producing the BLA - OVI absorption in each of these cases could also possibly be understood using the nonequilibrium *hybrid* scenarios similar to the $z = 0.01027$ system described in this paper.

7.1. Summary

We have presented HST/COS high S/N observations of the $z(\text{OVI}) = 0.01027$ broad $\text{Ly}\alpha$ - OVI absorber in the spectrum of the Seyfert 1 galaxy Mrk 290. The absorber was discovered by Wakker & Savage (2009) from the identification of the OVI $\lambda 1032 \text{ \AA}$ line in the *FUSE* spectrum, which did not cover $\text{Ly}\alpha$. The main conclusions from our analysis are as follows :

(1) The COS high S/N G130M spectra show $\text{Ly}\alpha$ absorption associated with the $z(\text{OVI}) = 0.01027$ system. In addition, the combined COS G130M and G160M spectra cover lines from C II, Si II, Fe II, C III, Si III, Si IV, C IV and NV. All these ions are non-detections at the 3σ significance level. The OVI $\lambda\lambda 1031, 1037 \text{ \AA}$ lines are detected in the *FUSE* spectrum of Mrk 290.

(2) The $\text{Ly}\alpha$ shows absorption from two distinct components with $\log [N(\text{H I}) \text{ cm}^{-2}] = 14.35 \pm 0.01$, $b(\text{H I}) = 55 \pm 1 \text{ km s}^{-1}$, and $\log [N(\text{H I}) \text{ cm}^{-2}] = 14.07 \pm 0.01$, $b(\text{H I}) = 33 \pm 1 \text{ km s}^{-1}$ at velocities of $v = 2 \pm 5 \text{ km s}^{-1}$ and $v = 116 \pm 5 \text{ km s}^{-1}$ in the rest frame of the absorber. The $v = 2 \text{ km s}^{-1}$ component is by definition a BLA ($b > 40 \text{ km s}^{-1}$). The BLA absorption profile is highly symmetric and Gaussian, with a dominant thermal broadening component of $b_{\text{thermal}}/b_{\text{total}} \sim 0.5$.

(3) For the OVI, we measure line parameters of $\log [N(\text{OVI}) \text{ cm}^{-2}] = 13.80 \pm 0.05$, $b(\text{OVI}) = 34 \pm 5 \text{ km s}^{-1}$. The OVI $\lambda 1032 \text{ \AA}$ absorption is kinematically well aligned with the BLA, with a $\Delta v = v(\text{OVI}) - v(\text{Ly}\alpha) = -1 \pm 5 \text{ km s}^{-1}$. The different b -values for the BLA and OVI yields a kinetic temperature of $T = 1.4 \times 10^5 \text{ K}$ in the gas, implying that $\sim 50\%$ of the line broadening is thermal.

(4) The Mrk 290 sight line is intercepting a filament of galaxies at the location of the absorber. The megaparsec scale filamentary structure extends over $\sim 20 \text{ deg}$ in the

sky, with several luminous galaxies distributed within $\sim 1.5 h^{-1} \text{ Mpc}$ projected distance and $\Delta v = 1000 \text{ km s}^{-1}$ systemic velocity of the absorber. Among the galaxies, NGC 5987 is closest in impact parameter and velocity to the absorber. The galaxy is at a projected distance of $\rho = 424 h^{-1} \text{ kpc}$, at a systemic velocity of $|v| = 71 \text{ km s}^{-1}$ and oriented almost edge on with respect to the absorber. NGC 5987 is an Sb galaxy with an isophotal major diameter of $D_{25} = 51.7 \text{ kpc}$ and a B-band magnitude of $B = 12.72$ or $M_B - 5 \log h = -20.4$ corresponding to a luminosity of $\sim 2L^*$.

(5) Photoionization models do not provide physically realistic solutions for the ion ratios in the absorber. The model that best satisfies the line strength ratios requires the absorbing region to have a very low density of $n_{\text{H}} \lesssim 6 \times 10^{-6} \text{ cm}^{-3}$ and a path length of $L \gtrsim 3 \text{ Mpc}$. Such a large path length would result in Hubble flow broadening that is inconsistent with the observed velocity widths of H I, OVI, and their relatively simple and symmetric line profiles. Also the temperature prediction of $T \sim 0.6 \times 10^5 \text{ K}$ made by the best-fit photoionization model is significantly smaller than $T = 1.4 \times 10^5 \text{ K}$ obtained from the different BLA and OVI b -values.

(6) We find that the existing simple CIE models, noneq-CI cooling models as well as *hybrid* models of CIE and photoionization do not yield a satisfactory explanation for the ionization in the absorber. The predictions most consistent with the observations are from *hybrid* models of non-equilibrium collisional ionization and photoionization. Within the frame work of such a *hybrid* model, a metallicity of $-1.7 \pm 0.1 \text{ dex}$ and ionization parameter of $-1.4 \pm 0.1 \text{ dex}$ corresponding to a density of $n_{\text{H}} \sim 4 \times 10^{-5} \text{ cm}^{-3}$, and a baryonic column density of $N(\text{H}) \sim 4 \times 10^{19} \text{ cm}^{-2}$ fits the observed ionic ratios at $T = 1.4 \times 10^5 \text{ K}$. The value of the derived metallicity is highly dependent on the validity of the model. In this model, the BLA is tracing collisionally ionized hydrogen whereas much of the OVI is produced via photoionization.

(7) The galaxy filament and the detection of low-metallicity gas at $T = 1.4 \times 10^5 \text{ K}$ suggest that this absorber is tracing a shock-heated gaseous structure, consistent with a few different scenarios including an overdense region of the WHIM in the galaxy filament, or the ionized gaseous halo of a galaxy in the filament (e.g. NGC 5987). In the latter scenario, the absorber could be a fragment accreted by the galaxy from the WHIM.

(8) Identifying metals associated with BLAs is an efficient way to detect plasma of $10^{4.5} \lesssim T \lesssim 10^6 \text{ K}$ associated with structures that could account for the remaining $\sim 50\%$ fraction of the missing baryons in the low- z universe. More such detections of thermally broadened $\text{Ly}\alpha$ and OVI absorbers are going to be facilitated by the high sensitivity spectroscopy capabilities of the *Cosmic Origins Spectrographs* in the far-UV.

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TABLE 1
COS OBSERVATIONS OF MRK 290

HST ID	Grating	Central Wavelength (Å)	Wavelength Range (Å)	Exposure Duration (sec)
LB4Q02010	G130M	1291	1136 – 1430	964
LB4Q02020	G130M	1300	1146 – 1440	964
LB4Q02030	G130M	1309	1156 – 1449	964
LB4Q02040	G130M	1318	1165 – 1459	964
LB4Q02050	G160M	1589	1400 – 1761	1200
LB4Q02060	G160M	1600	1412 – 1773	1200
LB4Q02070	G160M	1611	1424 – 1784	1200
LB4Q02080	G160M	1623	1436 – 1796	1200

NOTE. — All observations were carried out on 28 October 2009 as part of cycle 17 COS-GTO proposal ID 1152 (P.I. James Green).

TABLE 2
LINE MEASUREMENTS FOR THE $z(\text{OVI}) = 0.01027$ ABSORBER

Line (Å)	Instrument	W_r (mÅ)	v (km s ⁻¹)	$\log [N \text{ (cm}^{-2}\text{)}]$ dex	b (km s ⁻¹)	$[-v, +v]$ (km s ⁻¹)	Method
Ly α	COS/G130M	...	2 ± 5	14.35 ± 0.01	55 ± 1	...	Fit
		...	116 ± 5	14.07 ± 0.01	33 ± 1	...	
Ly α		522 ± 10	...	14.30 ± 0.01	64 ± 5	$[-220, 70]$	AOD
		293 ± 10	...	14.01 ± 0.01	41 ± 4	$[70, 250]$	
Ly α		808 ± 10	...	14.48 ± 0.01	98 ± 4	$[-220, 250]$	AOD
Ly γ	FUSE	< 190	...	< 15.2	...	$[-220, 250]$	3σ
$v(\text{Ly}\alpha) \sim 2$ COMPONENT							
O VI 1032	FUSE	...	1 ± 5	13.80 ± 0.05	29 ± 3	...	Fit
O VI 1032	FUSE	74 ± 13	-4 ± 4	13.82 ± 0.08	44 ± 6	$[-75, 75]$	AOD
O VI 1032	FUSE	49 ± 11	...	13.63 ± 0.08	34 ± 4	...	AOD [†]
N V 1239	COS/G130M	< 35	...	< 12.9	...	$[-75, 75]$	3σ
N V 1243	COS/G130M	< 33	...	< 13.0	...	$[-75, 75]$	3σ
C IV 1548	COS/G130M	< 48	...	< 12.9	...	$[-75, 75]$	3σ
C IV 1551	COS/G130M	< 43	...	< 13.1	...	$[-75, 75]$	3σ
Si IV 1394	COS/G130M	< 42	...	< 12.4	...	$[-75, 75]$	3σ
C III 977	COS/G130M	< 114	...	< 13.3	...	$[-75, 75]$	3σ
Si III 1207	COS/G130M	< 42	...	< 12.1	...	$[-20, 75]$	3σ
C II 1036	COS/G130M	< 36	...	< 13.4	...	$[-75, 75]$	3σ
C II 1335	COS/G130M	< 43	...	< 13.1	...	$[-75, 75]$	3σ
Si II 1190	COS/G130M	< 36	...	< 12.5	...	$[-75, 75]$	3σ
Fe II 1145	COS/G130M	< 33	...	< 13.5	...	$[-75, 75]$	3σ
$v(\text{Ly}\alpha) \sim 116$ COMPONENT							
O VI 1038	FUSE	< 31	...	< 13.5	...	$[75, 150]$	AOD
N V 1239	COS/G130M	< 33	...	< 12.4	...	$[75, 150]$	3σ
C IV 1548	COS/G130M	< 39	...	< 12.7	...	$[75, 150]$	3σ
C III 977	COS/G130M	26 ± 27	103 ± 7	< 13.2	...	$[75, 150]$	3σ
C III 977	COS/G130M	57 ± 22	103 ± 7	13.11 ± 0.22	...	$[80, 130]$	3σ
Si III 1207	COS/G130M	< 37	...	< 12.0	...	$[75, 150]$	3σ
C II 1335	COS/G130M	< 40	...	< 13.0	...	$[75, 150]$	3σ

NOTE. — Column 3 lists the rest-frame equivalent width, column 4 gives the velocity centroid of line components derived from profile fitting, column 5 lists the column density, column 6 is the Doppler parameter and column 7 the integration velocity range to derive the respective W_r , N and b .

All lines other than Ly α and O VI are non-detections at the 3σ significance level. For the two components in Ly α , different velocity ranges were chosen for measurements of the metal ions. The O VI $\lambda\lambda 1031, 1037$ Å, Ly γ and C III $\lambda 977$ Å are from *FUSE*.

For Ly α and O VI $\lambda 1032$ Å, we list the line parameters derived from fitting a Voigt function to the profile and also using the apparent optical depth (AOD) method of Sembach & Savage (1992). The profile fit was performed using the Fitzpatrick & Spitzer (1997) routine with the fit model for Ly α convolved with the COS/G130M line spread function described in Ghavamian et al. (2009). The velocity errors listed are dominated by the velocity calibration errors of $\sim \pm 5 \text{ km s}^{-1}$ for both COS and *FUSE*. For the *FUSE* O VI $\lambda 1032$ Å line a Gaussian kernel of $\text{FWHM} = 25 \text{ km s}^{-1}$ was convolved with the Voigt model to fit the observed profile.

[†] The measurement from Wakker & Savage (2009).

TABLE 3
BROAD Ly α ABSORBERS WITH METALS

Target	z	$\log [N(\text{H I}) \text{ cm}^{-2}]$ (dex)	$b(\text{H I})$ (km s $^{-1}$)	$\log [N(\text{O VI}) \text{ cm}^{-2}]$ (dex)	$b(\text{O VI})$ (km s $^{-1}$)	$ \Delta v $ (km s $^{-1}$)	T (K)	Reference
3C 273	0.00337	14.27 ± 0.08	53 ± 3	13.24 ± 0.21	34 ± 6	2	1.1×10^5	1
MRC 2251-178	0.00755	13.43 ± 0.05	64 ± 4	13.87 ± 0.26	35 ± 5	8	1.9×10^5	1
MRK 876	0.00312	14.27 ± 0.08	77 ± 2	13.18 ± 0.21	29 ± 3	9	3.3×10^5	1
PG 1444+407	0.22032	13.65 ± 0.05	86 ± 15	13.94 ± 0.07	36 ± 8	3	3.9×10^5	2
H 1821+643	0.12147	13.78 ± 0.17	85^{+37}_{-26}	14.02 ± 0.07	59^{+30}_{-20}	10	2.4×10^5	2
H 1821 + 643	0.26656	13.63 ± 0.02	44 ± 2	13.63 ± 0.04	25 ± 3	6	0.8×10^5	2
PG 1259+593	0.31978	13.98 ± 0.06	74 ± 9	13.49 ± 0.07	19 ± 4	10	3.3×10^5	3
$\log [N(\text{C III}) \text{ cm}^{-2}] \quad b(\text{C III})$								
H 1821 + 643	0.22638	13.50 ± 0.01	51 ± 2	12.56 ± 0.05	28 ± 4	1	1.2×10^5	4

NOTE. — A listing of known BLAs with metals with absorption aligned within $|\Delta v| = 15 \text{ km s}^{-1}$ in velocity. The references are (1) Wakker & Savage (2009) (2) Tripp et al. (2008) (3) Richter et al. (2004) (4) Narayanan, Savage & Wakker (2010). Column 1 is the sight line, column 2 the redshift of the absorber, column 3 and 5 are the column densities of H I and metal ion, column 4 and 6 are their respective b -values, column 7 is the difference in velocity between the Ly α and the metal-line, column 8 is the temperature derived using the b -values for H I and the metal ion assuming that the two species live in the same gas phase and the listed b is a quadrature sum of thermal and non-thermal b components.

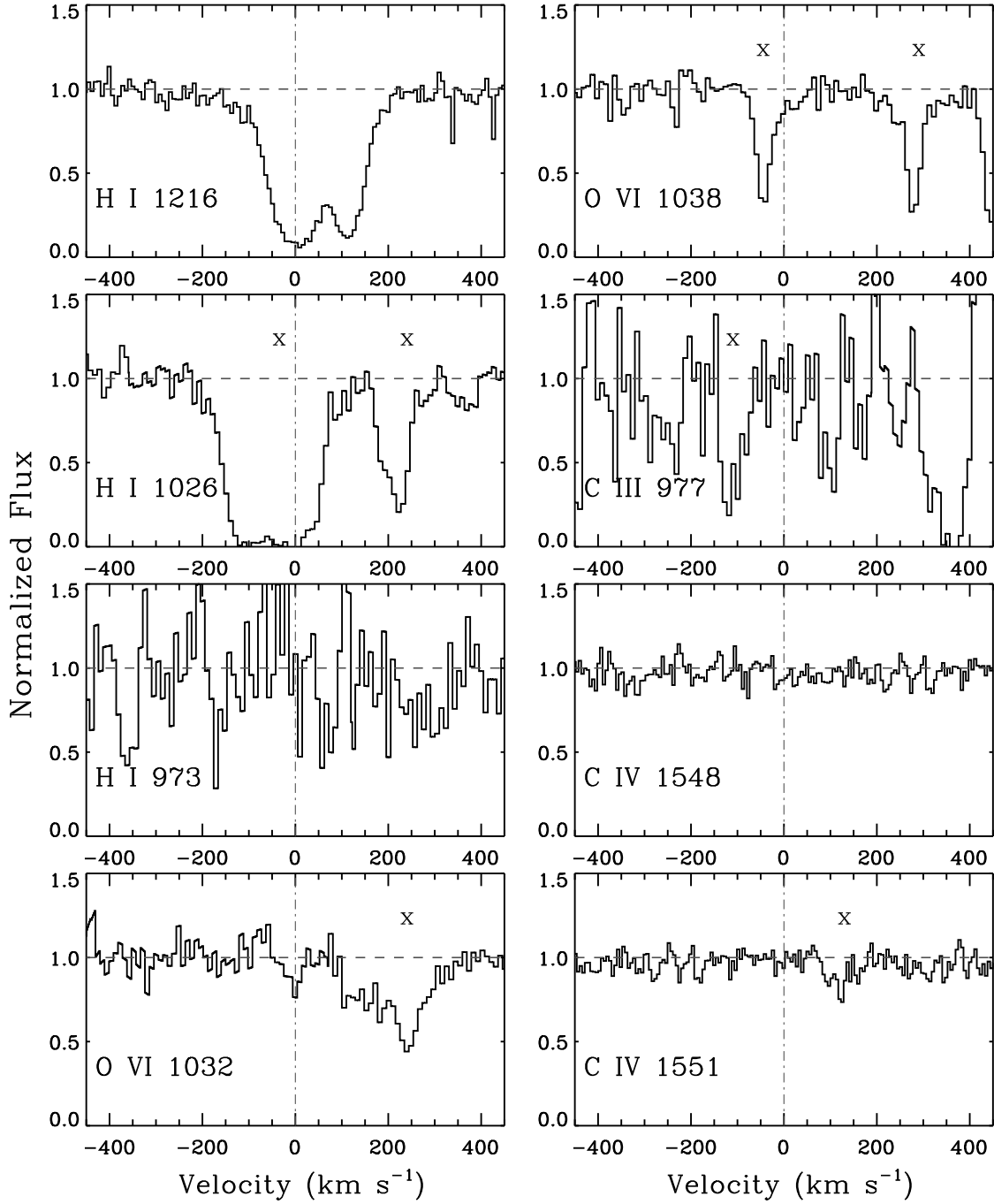


Fig. 1a. — Continuum normalized spectrum of Mrk 290 showing absorption profiles of Ly α , and OVI and the wavelength regions of other prominent lines in the $z(\text{OVI}) = 0.01027$ rest-frame of the absorption system. The OVI $\lambda\lambda 1031, 1037$ Å, Ly γ and C II $\lambda 1036$ Å, Ly β and Ly γ are lower S/N *FUSE* spectra (Wakker & Savage 2009) and the rest are *HST*/COS G130M and G160M observations. Every other metal line, except OVI, is a non-detection at the 3σ significance level. The line measurements are listed in Table 2. Ly β and OVI $\lambda 1038$ Å are strongly contaminated by Galactic C II $\lambda 1036$ Å and Ar I $\lambda 1048$ Å respectively. The wavelength region corresponding to the Ly γ does not suggest any significant unresolved saturation in H I. Features which are not part of the absorption system are labeled with "x" in the respective panels.

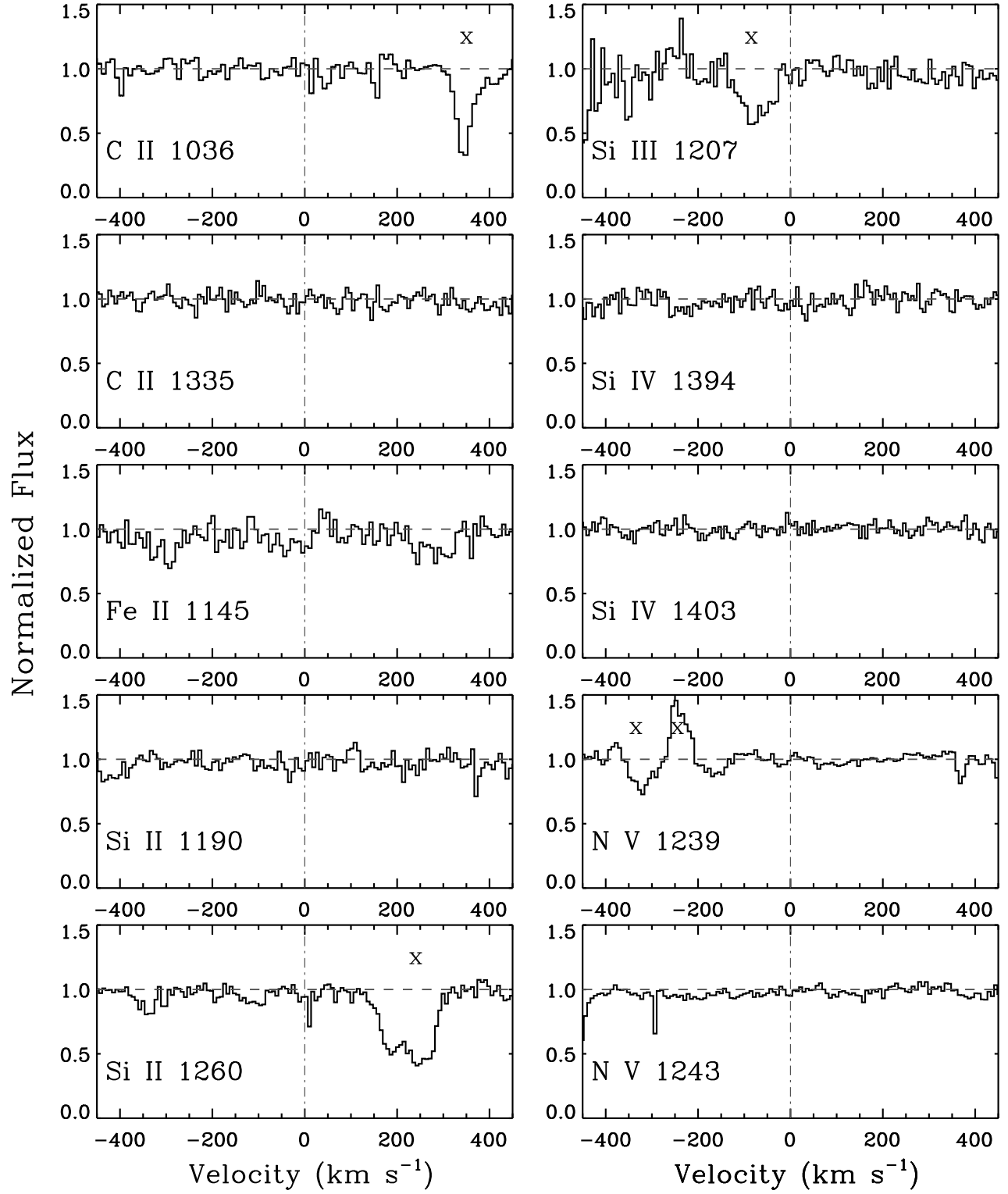


Fig. 1b. — Continuation of Figure 1a.

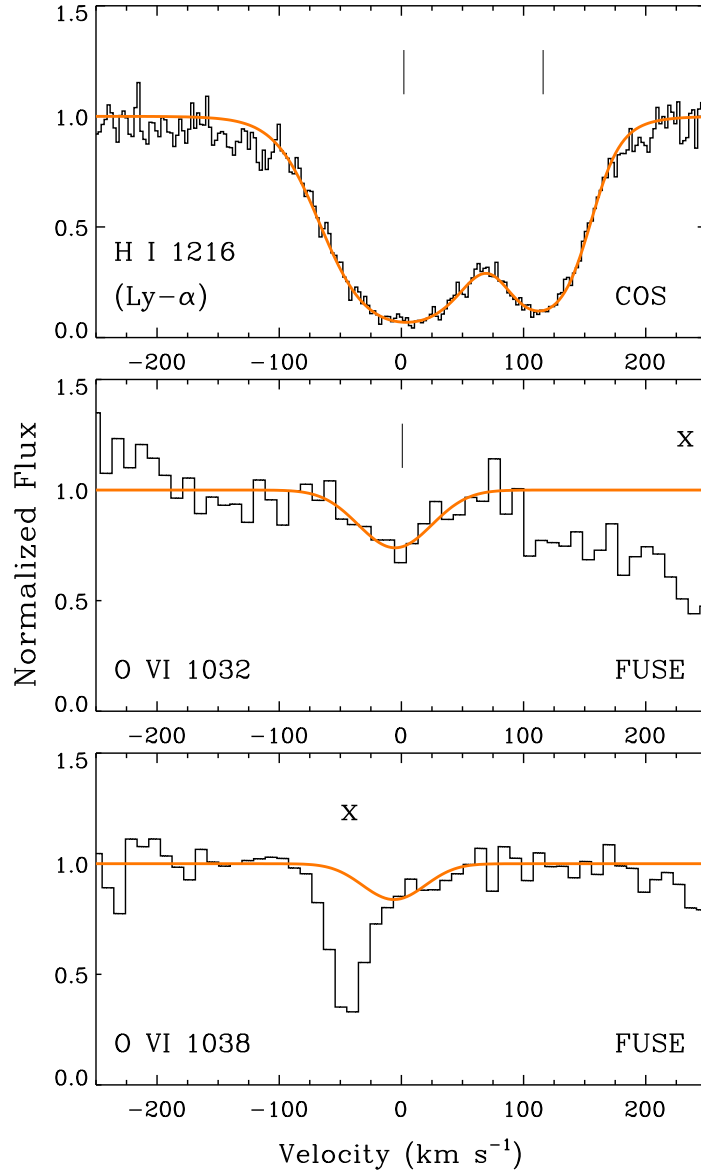


FIG. 2.— *Top Panel* : Shows the two component Voigt profile model superimposed on the Ly α absorption. The vertical axis is continuum normalized flux and the horizontal axis is velocity in the rest-frame of the absorber, where $v = 0 \text{ km s}^{-1}$ corresponds to $z(\text{OVI}) = 0.01027$. Features which are not part of the absorption system are marked using "x" in the respective panels. The tick marks above the Ly α profile show the centroids of the two absorbing components at $v = 2 \pm 5 \text{ km s}^{-1}$ and $v = 116 \pm 5 \text{ km s}^{-1}$ respectively. The model Voigt profiles were convolved with the COS/G130M line spread function given by Ghavamian et al. (2009). *Middle & Bottom Panels* : Show the fit to the OVI $\lambda 1032 \text{ \AA}$ and the contaminated OVI $\lambda 1038 \text{ \AA}$ lines respectively. A Gaussian of FWHM = 25 km s^{-1} corresponding to the resolution of FUSE was used to derive the model fit. The fit parameters are given in Table 2.

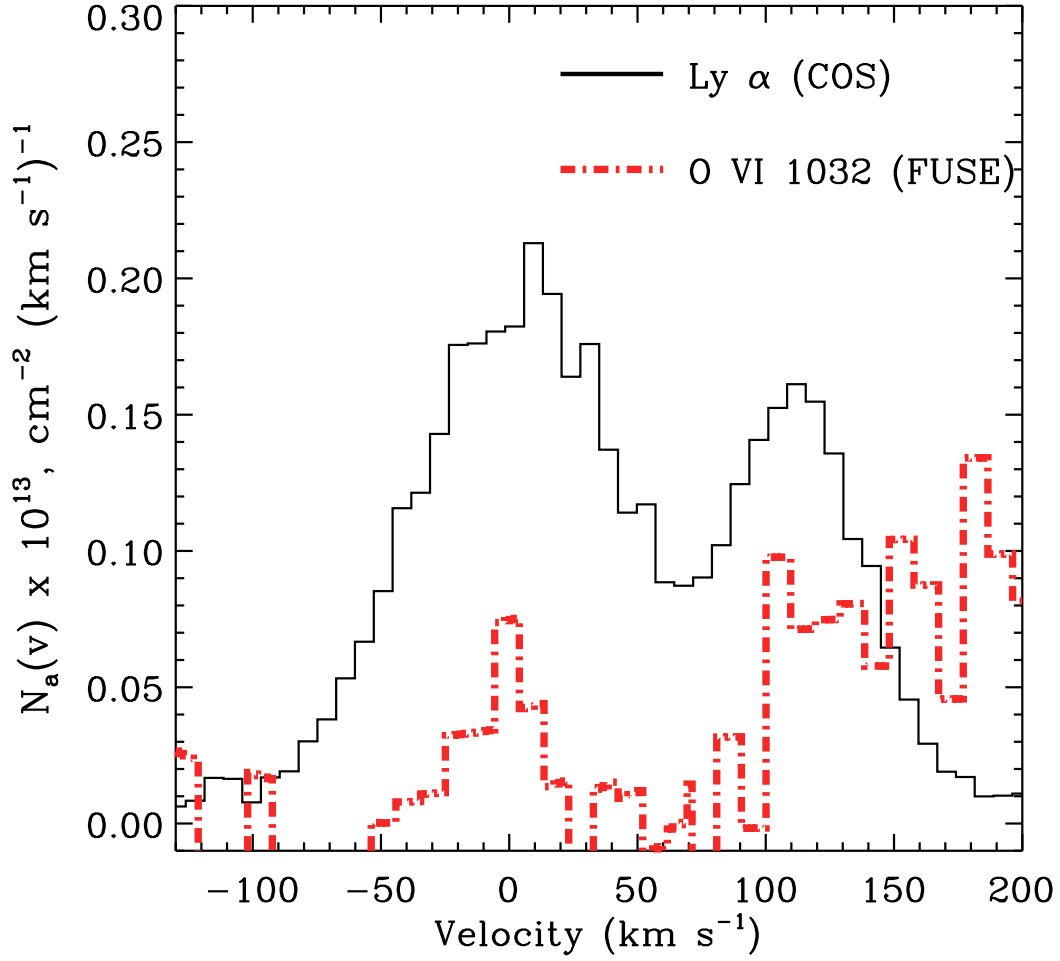


FIG. 3.— Apparent column density ($N_a(v)$) profile plots for $\text{Ly}\alpha$ and $\text{OVI } \lambda 1032 \text{ \AA}$. The $\text{OVI } \lambda 1032 \text{ \AA}$ absorption at $v = 1 \pm 5 \text{ km s}^{-1}$ is kinematically well-aligned with the BLA component of the H I absorption at $v = 2 \pm 5 \text{ km s}^{-1}$, suggesting a single-phase origin for H I and OVI . The $\text{Ly}\alpha$ and OVI are binned to approximately the same resolution of $\sim 9 \text{ km s}^{-1}$ velocity bins. The $N_a(v)$ for $\text{OVI } \lambda 1032$ is contaminated at $v \geq 100 \text{ km s}^{-1}$ by a feature unrelated to this system.

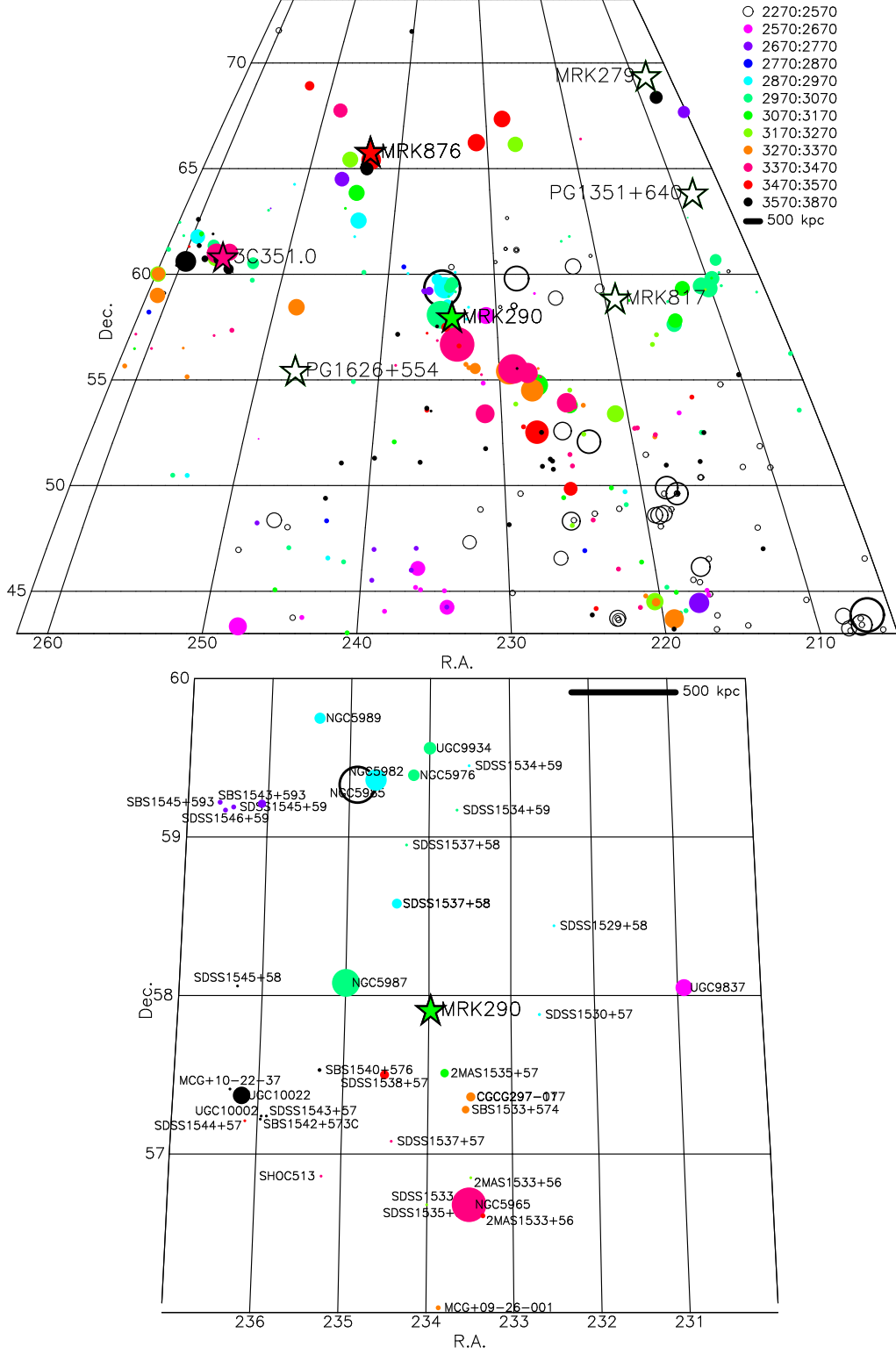


Fig. 4a. — *Top Panel* : Galaxies in the foreground field of Mrk 290 within a systemic velocity of $v \sim 1000 \text{ km s}^{-1}$ and projected distance of $\sim 4 h^{-1} \text{ Mpc}$ from the BLA-OVI absorber. The *circles* represent galaxies and the *stars* represent sight-lines towards AGNs in the field. The distribution pattern of the galaxies shows that the sight line to Mrk 290 is intercepting a megaparsec scale filament of galaxies at the location of the BLA - OVI absorber. Those AGNs with an intervening absorption system identified at the redshift of the galaxy filament are plotted in *filled star* symbols. The size of the *circles* are proportional to the size of the galaxies given by their D_{25} major diameter. The color coding of the symbols represent various heliocentric velocity ranges. The information on galaxies and AGNs in this region of the sky were retrieved from NED. *Bottom Panel* : A zoom-in view of the $\sim 1 h^{-1} \text{ kpc}$ region centered around Mrk 290. The galaxy IDs are labeled. Among the extended sources, NGC 5987 is closest in impact parameter ($\rho = 424 h^{-1} \text{ kpc}$) and velocity ($|\Delta v| = 73 \text{ km s}^{-1}$) to the BLA - OVI absorber.

Fig. 4b.— A Digital Sky Server (DSS) image of NGC 5987 with the line of sight to Mrk 290 marked. Figure not included in the astro-ph version due to large file size.

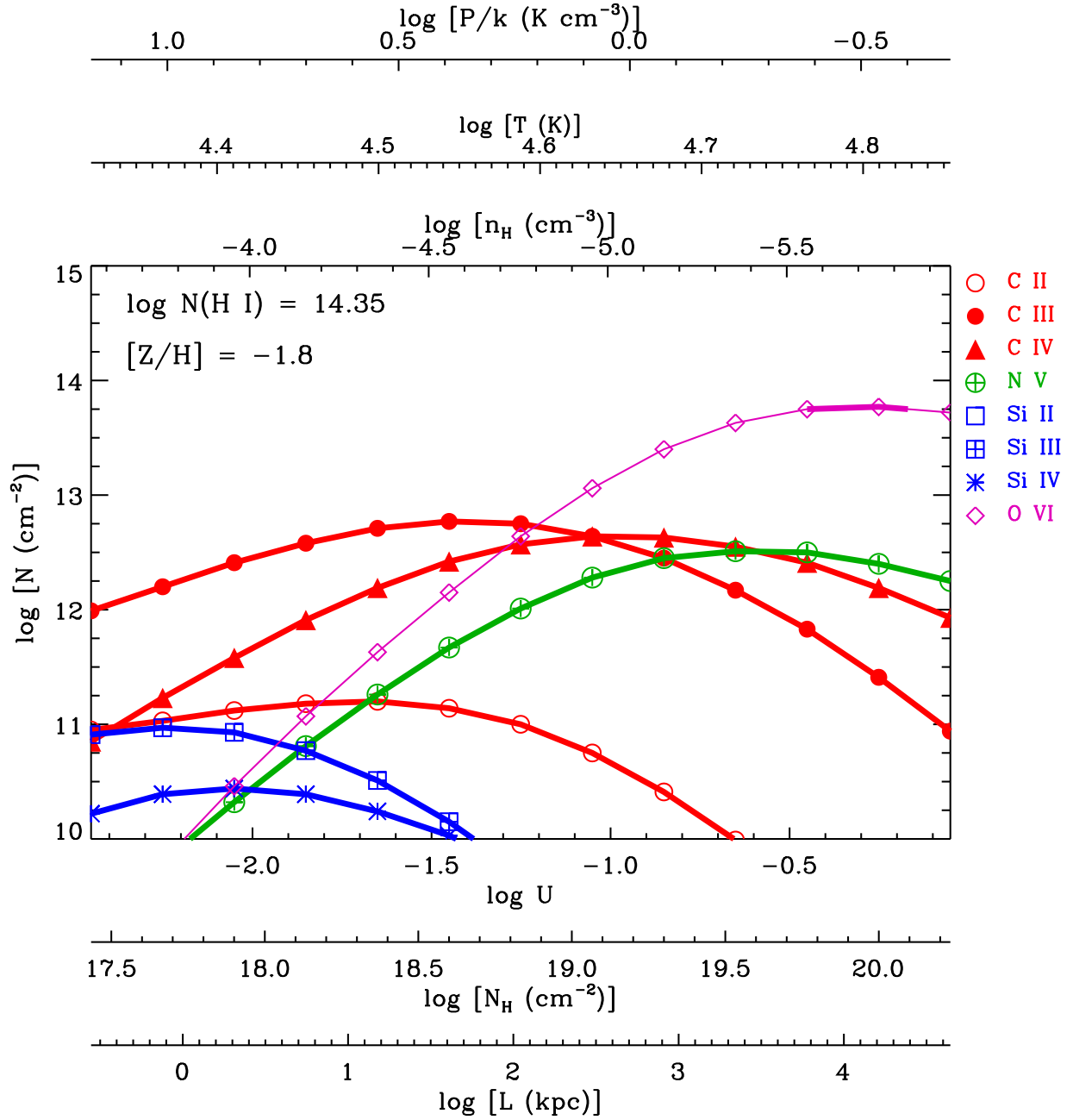


FIG. 5.— Photoionization model predictions of column densities for the BLA - OVI absorber. The $N(\text{H I}) = 10^{14.35} \text{ cm}^{-2}$ is the column density for the $v = 2 \text{ km s}^{-1}$ BLA component of the H I absorption. The ionizing background is from the Haardt & Madau (2001) model for $z = 0.01027$ and includes UV photons from quasars and young star forming galaxies. The acceptable range of column densities for each ion, based on measurement, are highlighted in the photoionization curves using *thick* lines. Except for OVI, all other ions are non-detections at the 3σ significance level. The column density measurements are therefore upper limits for those ions. The multiple X-axes are the ionization parameter ($\log U$), total hydrogen column density ($N(\text{H})$, cm^{-2}), hydrogen number density (n_{H} , cm^{-3}), temperature (T , K), gas pressure (p/K , $\text{cm}^{-3} \text{ K}$, and size of the absorbing region ($L = N(\text{H})/n_{\text{H}}$, cm). The models assume solar relative elemental abundances of Asplund et al. (2009). The metallicity for which the models are able to recover the measured $N(\text{OVI})$ is -1.8 dex of solar.

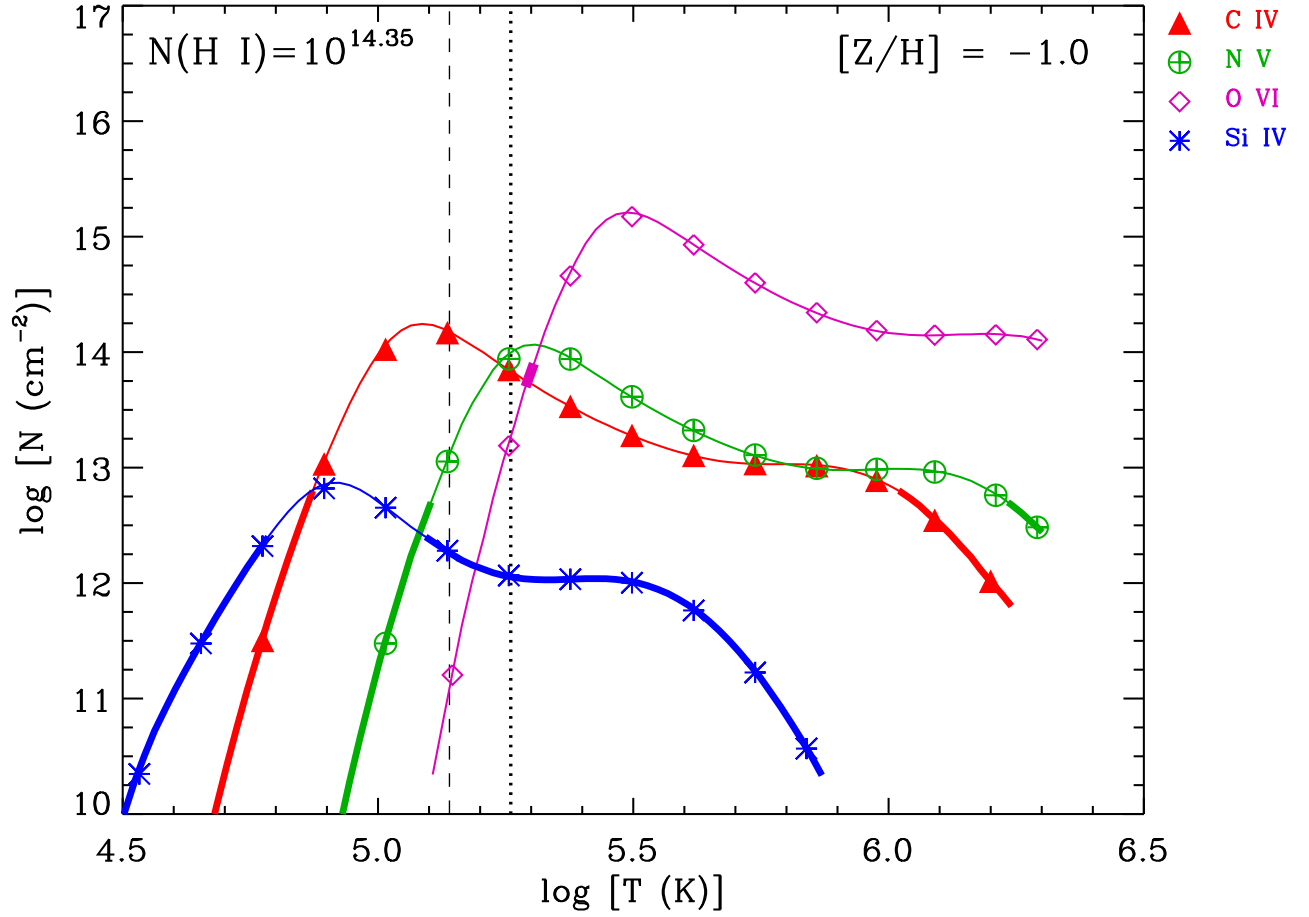


FIG. 6.— Collisional ionization equilibrium curves for the high ions in the BLA - OVI absorber based on the models of Gnat & Sternberg (2007). The H I column density in these models are set to the measured value of $10^{14.35} \text{ cm}^{-2}$ (see table 2). We use the updated solar elemental abundances of Asplund et al. (2009). Except OVI, the other high ions are non-detections at $\geq 3 \sigma$ significance. Their column densities are therefore upper limits. The region of the OVI CIE curve where the measured $N(\text{OVI}) \pm \sigma[N(\text{OVI})]$ is recovered is drawn in *thick*. Similarly, the regions of the C IV, N V and Si IV CIE curves that are consistent with the measured upper limits of their column densities are drawn using *thick* lines. The vertical *dashed* line marks $T = 1.4 \times 10^5 \text{ K}$, the temperature given by the combined BLA and OVI b -values. The vertical *dotted* line corresponds to the temperature upper limit of $T = 1.8 \times 10^5 \text{ K}$ if the entire $b(\text{H I}) = 55 \text{ km s}^{-1}$ line width for the BLA is thermal.

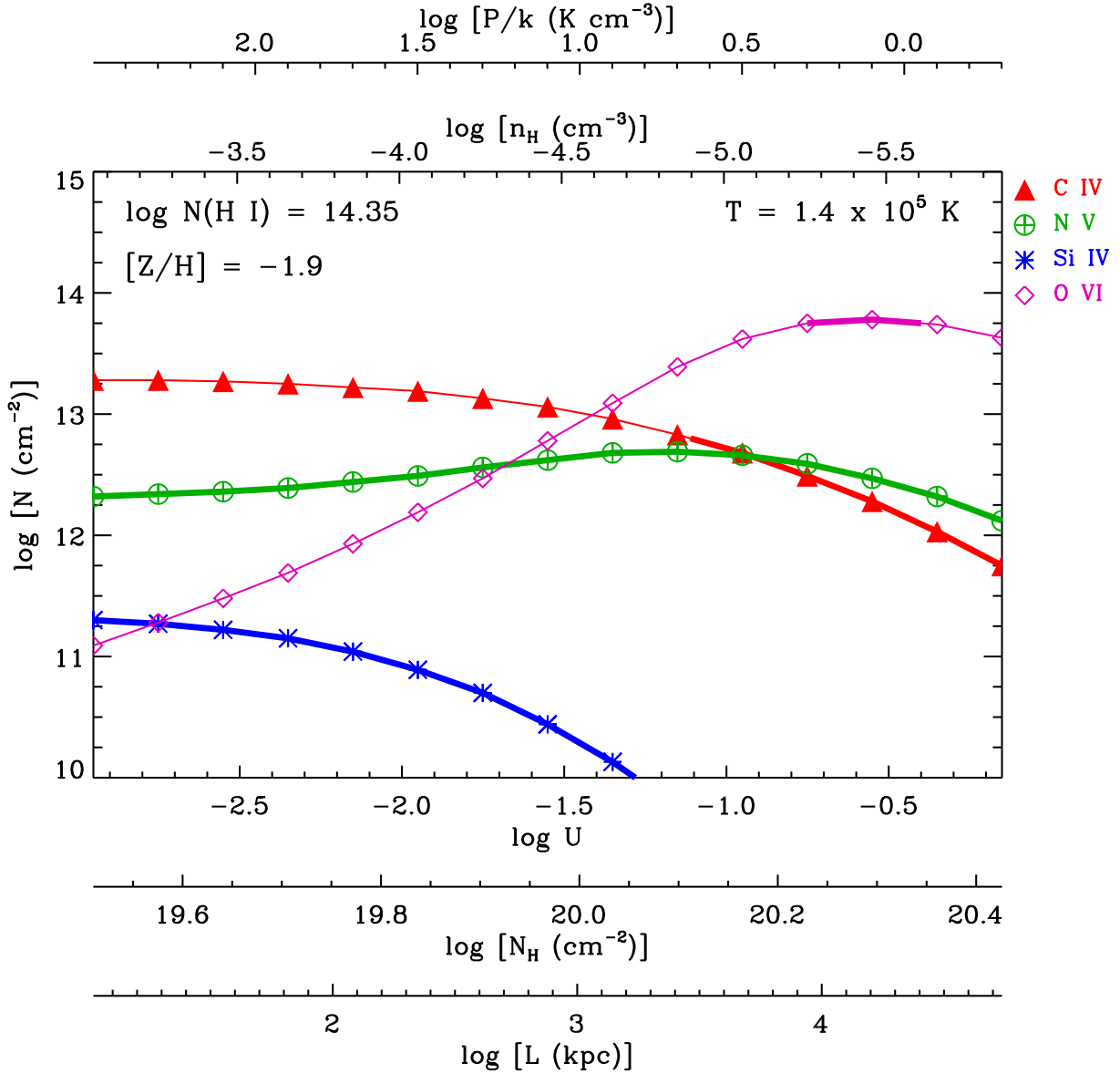


FIG. 7.— Column density predictions from *hybrid* models where both collisional and photoionization reactions are allowed to contribute to ion fractions. The models were computed using Cloudy [ver.C08.00, Ferland et al. (1998)]. The temperature in the models is fixed to $T = 1.4 \times 10^5 \text{ K}$ derived for the BLA - OVI phase, and the $N(\text{H I}) = 10^{14.35} \text{ cm}^{-2}$ measured from profile fit to Ly α (see table 2). The acceptable column density ranges based on measurements for the various ions are drawn in *thick*. The observed $N(\text{OVI})$ is recovered at $\log U \sim -0.6$ for a metallicity of -1.9 dex, assuming solar elemental abundances of Asplund et al. (2009).

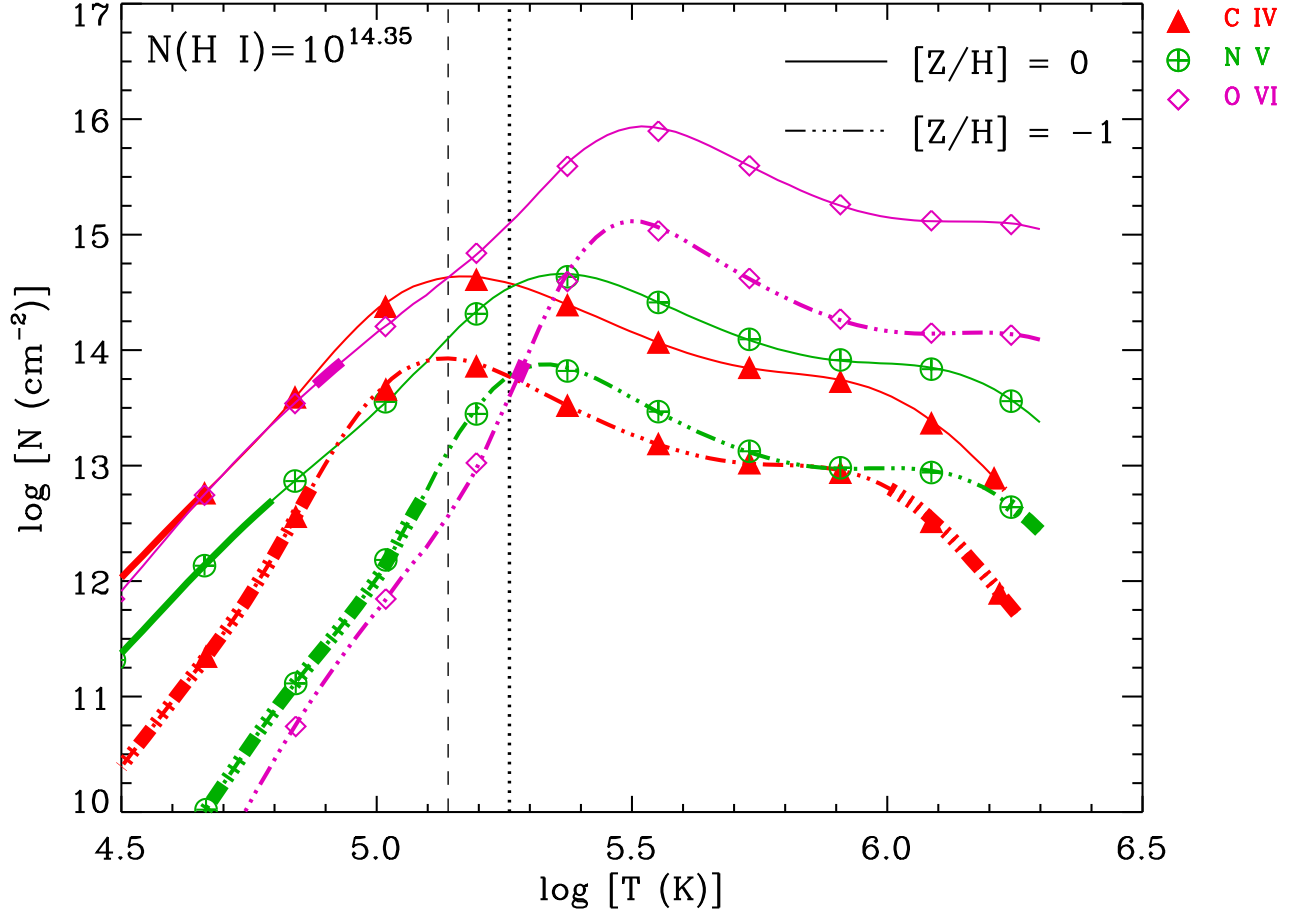


FIG. 8.— Nonequilibrium collisional ionization models of Gnat & Sternberg (2007) for 1/10 solar (*dash - dot* line) and solar (*solid* line) metallicities and updated solar relative elemental abundances of Asplund et al. (2009). The acceptable column density ranges for the various ions are drawn in *thick* on each ion curve. Due to rapid rate of cooling, which is higher for the higher gas metallicities, the ion fractions of O VI (and other high ions) are greater than the predictions of CIE in the temperature range $10^{4.5} \leq T \leq 10^{5.5}$ K. The vertical *dashed* line marks $T = 1.4 \times 10^5$ K, the temperature given by the combined BLA and O VI b -values. The vertical *dotted* line corresponds to the temperature upper limit of $T = 1.8 \times 10^5$ K if the entire $b(\text{H I}) = 55 \text{ km s}^{-1}$ line width for the BLA is thermal.

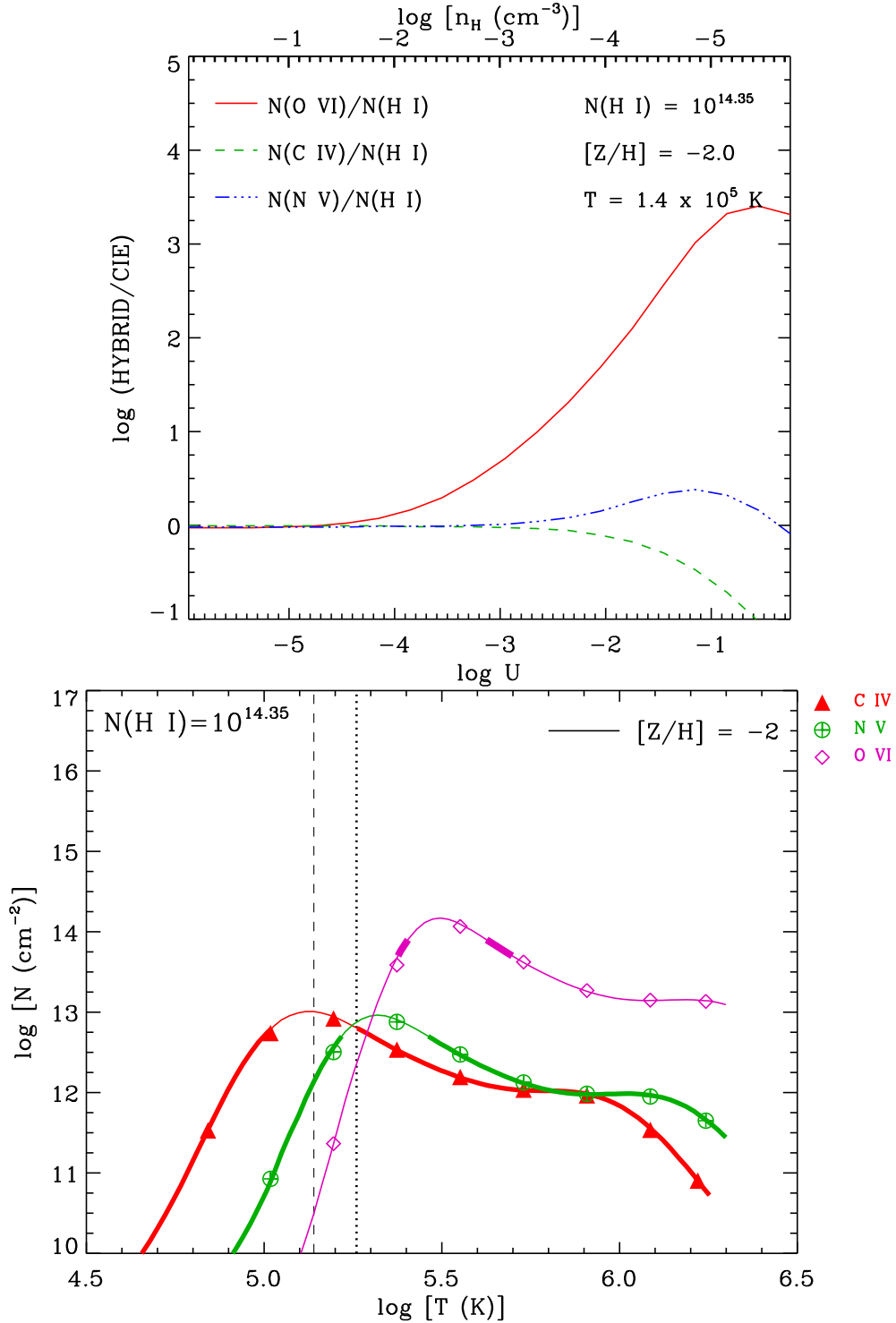


FIG. 9.— *Top Panel* : CIE - photoionization *hybrid* models computed using Cloudy [ver.C08.00, Ferland et al. (1998)] for solar metallicity. The Y-axis shows the ion ratio predictions in the *hybrid* models normalized to their corresponding value for CIE. *Bottom Panel* : Column density predictions made by noneq-CI models of Gnat & Sternberg(2007) for a radiatively cooling gas at different temperatures and a metallicity of 1/100 solar. The vertical *dashed* line marks $T = 1.4 \times 10^5 \text{ K}$, the temperature given by the combined BLA and OVI *b*-values. The vertical *dotted* line corresponds to the temperature upper limit of $T = 1.8 \times 10^5 \text{ K}$ if the entire $b(\text{H I}) = 55 \text{ km s}^{-1}$ line width for the BLA is thermal. To approximately allow for photoionization, the model predictions in the *bottom panel* at $T = 1.4 \times 10^5 \text{ K}$ need to be corrected by the values for each ion given in the *left panel* for the appropriate value of $\log U$.